

## Seeds of Life in Space (SOLIS)

### VII. Discovery of a cold dense methanol blob toward the L1521F VeLLO system<sup>★</sup>

C. Favre<sup>1</sup>, C. Vastel<sup>2</sup>, I. Jimenez-Serra<sup>3</sup>, D. Quénard<sup>4</sup>, P. Caselli<sup>5</sup>, C. Ceccarelli<sup>1</sup>, A. Chacón-Tanarro<sup>6</sup>, F. Fontani<sup>7,5</sup>, J. Holdship<sup>8,9</sup>, Y. Oya<sup>10</sup>, A. Punanova<sup>11</sup>, N. Sakai<sup>12</sup>, S. Spezzano<sup>5</sup>, S. Yamamoto<sup>10</sup>, R. Neri<sup>13</sup>, A. López-Sepulcre<sup>1,13</sup>, F. Alves<sup>5</sup>, R. Bachiller<sup>6</sup>, N. Balucani<sup>14,7,1</sup>, E. Bianchi<sup>1</sup>, L. Bizzocchi<sup>5</sup>, C. Codella<sup>7,1</sup>, E. Caux<sup>2</sup>, M. De Simone<sup>1</sup>, J. Enrique Romero<sup>1,15</sup>, F. Dulieu<sup>16</sup>, S. Feng<sup>5</sup>, A. Jaber Al-Edhari<sup>1,17</sup>, B. Lefloch<sup>1</sup>, J. Ospina-Zamudio<sup>1</sup>, J. Pineda<sup>5</sup>, L. Podio<sup>7</sup>, A. Rimola<sup>15</sup>, D. Segura-Cox<sup>5</sup>, I. R. Sims<sup>18</sup>, V. Taquet<sup>7</sup>, L. Testi<sup>19,7</sup>, P. Theulé<sup>20</sup>, P. Ugliengo<sup>21</sup>, A. I. Vasyunin<sup>22,11</sup>, F. Vazart<sup>1</sup>, S. Viti<sup>8</sup>, and A. Witzel<sup>1</sup>

(Affiliations can be found after the references)

Received 11 December 2019 / Accepted 12 February 2020

#### ABSTRACT

**Aims.** The Seeds Of Life In Space IRAM/NOEMA large program aims at studying a set of crucial complex organic molecules in a sample of sources with a well-known physical structure that covers the various phases of solar-type star formation. One representative object of the transition from the prestellar core to the protostar phases has been observed toward the very low luminosity object (VeLLO) L1521F. This type of source is important to study to link prestellar cores and Class 0 sources and also to constrain the chemical evolution during the process of star formation.

**Methods.** Two frequency windows (81.6–82.6 GHz and 96.65–97.65 GHz) were used to observe the emission from several complex organics toward the L1521F VeLLO. These setups cover transitions of ketene ( $\text{H}_2\text{CCO}$ ), propyne ( $\text{CH}_3\text{CCH}$ ), formamide ( $\text{NH}_2\text{CHO}$ ), methoxy ( $\text{CH}_3\text{O}$ ), methanol ( $\text{CH}_3\text{OH}$ ), dimethyl ether ( $\text{CH}_3\text{OCH}_3$ ), and methyl formate ( $\text{HCOOCH}_3$ ).

**Results.** Only two transitions of methanol ( $\text{A}^+$ ,  $\text{E}_2$ ) have been detected in the narrow window centered at 96.7 GHz (with an upper limit on  $\text{E}_1$ ) in a very compact emission blob ( $\sim 7''$  corresponding to  $\sim 1000$  au) toward the northeast of the L1521F protostar. The CS 2–1 transition is also detected within the WideX bandwidth. Consistently with what has been found in prestellar cores, the methanol emission appears  $\sim 1000$  au away from the dust peak. The location of the methanol blob coincides with one of the filaments that have previously been reported in the literature. The excitation temperature of the gas inferred from methanol is  $(10 \pm 2)$  K, while the  $\text{H}_2$  gas density (estimated from the detected CS 2–1 emission and previous CS 5–4 ALMA observations) is a factor  $>25$  higher than the density in the surrounding environment ( $n(\text{H}_2) \geq 10^7 \text{ cm}^{-3}$ ).

**Conclusions.** Based on its compactness, low excitation temperature, and high gas density, we suggest that the methanol emission detected with NOEMA is (i) either a cold and dense shock-induced blob that formed recently ( $\leq$  a few hundred years) by infalling gas or (ii) a cold and dense fragment that may just have been formed as a result of the intense gas dynamics within the L1521F VeLLO system.

**Key words.** astrochemistry – line: identification – ISM: abundances – ISM: molecules – ISM: individual objects: L1521F

### 1. Introduction

About a decade ago, the *Herschel* satellite revealed that the interstellar matter (ISM) is organized in a complex network of filamentary structures or filaments (André et al. 2010). These filaments are believed to undergo gravitational fragmentation into multiple fragments that subsequently form dense and cold prestellar cores (see, e.g., André et al. 2019).

It is well established that low-mass protostars are born within prestellar cores. However, the transition between a prestellar core and a protostar (called first hydrostatic core phase, or FHSC) is poorly known. The stage of the FHSC starts when the density of the central object increases enough (through accretion) to become opaque to radiation, and lasts until its temperature reaches 2000 K, which forces the dissociation of  $\text{H}_2$

(Larson 1969; Masunaga et al. 1998; Masunaga & Inutsuka 2000). Because their lifetimes are short (0.5–50 kyr, Omukai 2007; Tomida et al. 2010; Commerçon et al. 2012), it is challenging to identify FHSCs.

Several observational studies have attempted to search and identify FHSCs (Dunham et al. 2008, 2011; Chen et al. 2010, 2012; Enoch et al. 2010; Schnee et al. 2012; Murillo & Lai 2013). Some of them have proposed that very low luminosity objects (or VeLLOs, with  $L_{\text{bol}} \leq 0.1 L_{\odot}$ ) might be FHSCs (e.g., Enoch et al. 2010). However, their true nature is still under debate. For example, Vorobyov et al. (2017) argued that the majority of VeLLOs should be in the evolved Class I protostars phase, where protostars have already grown in mass through so-called cold accretion (i.e., a phenomenon by which the accreting gas provides very low entropy to the protostar; Hosokawa et al. 2011). Interestingly, Tokuda et al. (2017) have found that the very low luminosity protostar in the L1521F system has a central stellar mass of  $\sim 0.2 M_{\odot}$ , and they suggested that this finding can

<sup>★</sup> Based on observations carried out under project number L15AA with the IRAM-NOEMA interferometer. IRAM is supported by INSU/CNRS (France), MPG (Germany), and IGN (Spain).

likely be explained by the cold accretion model. Therefore, studies of VeLLOs are not only important for understanding the earliest evolutionary stages in the formation process of low-mass stars, but they may also represent the missing link between the prestellar core phase and the Class 0 phase.

L1521F (also known as MC27, see Codella et al. 1997; Mizuno et al. 1994; Onishi et al. 1996, 1998, 1999, 2002) is one of the densest cores in the nearby ( $\sim 136$  pc; Maheswar et al. 2011) Taurus molecular cloud. It was originally classified as a starless core and was the subject of many studies; it shares many similarities with the prototypical prestellar core L1544. Crapsi et al. (2004) observed L1521F in dust emission at 1.2 mm and in two transitions each of  $\text{N}_2\text{H}^+$ ,  $\text{N}_2\text{D}^+$ ,  $\text{C}^{18}\text{O}$ , and  $\text{C}^{17}\text{O}$ . They measured a molecular hydrogen number density  $n(\text{H}_2) \sim 10^6 \text{ cm}^{-3}$  and a CO depletion factor, integrated along the line of sight, of  $f_D = 9.5 \times 10^{-5}/x_{\text{obs}}(\text{CO}) \sim 15$ , similar to that derived toward the prestellar core L1544. The  $\text{N}(\text{N}_2\text{D}^+)/\text{N}(\text{N}_2\text{H}^+)$  column density ratio is  $\sim 0.1$ , a factor of about 2 lower than that found in L1544. The  $\text{N}_2\text{H}^+$  and  $\text{N}_2\text{D}^+$  line widths in the core center are  $\sim 0.3 \text{ km s}^{-1}$ , significantly larger than in other more quiescent Taurus starless cores, but similar to those observed toward the center of L1544. From all this, Crapsi et al. (2004, 2005) concluded that L1521F is less evolved than L1544, but, in analogy with the latter core, it is approaching the “critical” state.

The view on the physical nature of the L1521F core changed through the high sensitivity of the *Spitzer* telescope, which detected a very low luminosity protostar ( $< 0.07 L_\odot$ ) in a very dense region ( $10^6 \text{ cm}^{-3}$ ; Bourke et al. 2006). L1521F-IRS is currently classified as a VeLLO (see, e.g., Young et al. 2004; Dunham et al. 2006; Lee et al. 2009).

Subsequent interferometric observations carried out with the SMA in  $^{12}\text{CO}$  (2–1) and 1.3 mm continuum emission spatially resolved a compact but poorly collimated molecular outflow associated with L1521F-IRS (Takahashi et al. 2013). This suggests that L1521F is at the earliest protostellar stage ( $< 10^4$  yr). In addition, higher angular resolution observations carried out with the IRAM-Plateau de Bure (PdBI) and ALMA showed that this source is split into a small cluster of cores, MMS-1, MMS-2, and MMS-3 (see Maury et al. 2010; Tokuda et al. 2014), where MMS-1 coincides with the location of the L1521F-IRS *Spitzer* source. The SMA observations of Takahashi et al. (2013) also unveiled another object in the region with evidence of compact CO blueshifted and redshifted components toward the northeast of L1521F-IRS, called L1521F-NE. However, no driving source has been detected in either millimeter continuum emission with PdBI/SMA or infrared emission with *Spitzer*, as confirmed by ALMA Cycle 1 observations (Tokuda et al. 2014, 2016). From this, Takahashi et al. (2013) derived a mass detection limit of  $10^{-4} M_\odot$  for L1521F-NE.

In this paper, we present new interferometric observations of L1521F carried out with the Northern Extended Millimetre Array (NOEMA) to investigate the molecular complexity of the identified VeLLO. This work is part of the NOEMA large program Seeds of Life in Space (SOLIS), which studies the formation of complex organic molecules across all stages of star formation (Ceccarelli et al. 2017). In Sect. 2 we present the details of the observations, the data reduction procedure, and the Gaussian fitting of the spectra. Section 3 presents the results of the Gaussian fitting, velocity gradients, rotational temperatures, and column density calculations of the  $\text{CH}_3\text{OH}$  and CS molecular lines detected toward L1521F with NOEMA. In Sect. 4 we discuss the results and possible origins of the methanol-rich blob that is located  $\sim 1000$  au away from the L1521F source.

## 2. Observations and data reduction

In the following subsections, we present the observations obtained through the IRAM-NOEMA interferometer and the IRAM-30 m telescope. A description of the data reduction and data merging is also given.

### 2.1. IRAM observations

#### 2.1.1. NOEMA observations

The IRAM-NOEMA observations were carried out in C and D configurations between September 2016 and January 2017 under average weather conditions (pwv = 1–10 mm) toward L1521F ( $\alpha_{2000} = 04^{\text{h}}28^{\text{m}}38.99^{\text{s}}$ ,  $\delta_{2000} = 26^\circ51'35.6''$ ). The rest frequencies were shifted with respect to the  $V_{\text{LSR}}$  of the source ( $\sim 6.4$ – $6.6 \text{ km s}^{-1}$ ). The primary beam size was  $52''$ , and the synthesized beam was  $2.72'' \times 2.37''$  at a position angle  $29^\circ$ . The data were obtained with the narrowband correlator with a spectral resolution of 39 kHz, corresponding to a velocity resolution of  $0.12 \text{ km s}^{-1}$ . The system temperatures were 50–110 K. The nearby sources 3C454.3 and J0438+300 were used as band-pass and gain (phase and amplitude) calibrators, respectively. The absolute flux calibration was performed by observing the quasar MWC34 (1.03 Jy).

Two methanol transitions were detected in the narrowband correlator:  $\text{E}_2 2_{1,2}-1_{1,1}$  (96.739362 GHz),  $\text{A}^+ 2_{0,2}-1_{0,1}$  (96.741375 GHz) in the northeast position of L1521F, while  $\text{E}_1 2_{0,2}-1_{0,1}$  (96.744550 GHz) was marginally detected. The emission is clearly extended (see Sect. 3.3). Along with the  $\text{E}_2 2_{1,2}-1_{1,1}$  (96.739362 GHz),  $\text{A}^+ 2_{0,2}-1_{0,1}$  (96.741375 GHz), and  $\text{E}_1 2_{0,2}-1_{0,1}$  (96.744550 GHz) methanol transitions, the dimethyl ether (E and A  $\text{CH}_3\text{OCH}_3 5_{5,1}-4_{4,0}$  at 95.85 GHz) and methyl formate (E- $\text{CH}_3\text{OCHO } 5_{4,1}-5_{3,3}$  at 96.94 GHz and A- $\text{CH}_3\text{OCHO } 17_{5,12}-17_{4,13}$  at 97.20 GHz) lines were observed within the same spectral setup with the narrowband correlator. Nonetheless, these molecular species (as well as the other targeted COMs, see Ceccarelli et al. 2017) are not detected in the map at high spectral resolution (rms  $\sim 3.8 \text{ mJy beam}^{-1}$ ). In addition, the spectral range of WideX was 95.85–99.45 GHz, and the CS (2–1) line at 97.98 GHz was detected with a spectral resolution of 1950 kHz ( $6.0 \text{ km s}^{-1}$ ), while the SO ( $2_3-1_2$ ) line at 99.30 GHz was not detected with an rms  $\sim 0.3$ – $0.4 \text{ mJy beam}^{-1}$  and a beam size of about  $3'' \times 2.6''$  (PA:  $24^\circ$ ).

In addition, observations at about 82 GHz in C and D configurations were also performed toward L1521F between September and November 2016 with eight antennas. However, only the continuum emission is detected in these datasets (see Sect. 3.1). None of the targeted lines (see Ceccarelli et al. 2017, for further details) was detected.

#### 2.1.2. IRAM-30 m observations

We here also use IRAM-30 m observations to recover the most extended emission. The single-dish observations were carried out in 2016 August under good weather conditions (pwv of about 1–2 mm). The on-the-fly maps were obtained with the EMIR 090 (3 mm band) heterodyne receiver in position-switching mode, using the FTS backend with a spectral resolution of 50 kHz; this corresponds to a velocity resolution of  $0.15 \text{ km s}^{-1}$  at the frequency of 96.74 GHz. The angular resolution was  $25.6''$ . The  $3' \times 3'$  maps were centered at the dust emission peak ( $\alpha_{2000} = 04^{\text{h}}28^{\text{m}}39.8^{\text{s}}$ ,  $\delta_{2000} = 26^\circ51'35''$ ). The pointing accuracy of the 30 m antenna was better than  $1''$ . The system

temperature was 157 K. A detailed description of the data will be given in an upcoming paper (Spezzano et al., in prep.).

## 2.2. Data reduction

The calibration, imaging, and cleaning of the NOEMA data were performed using the CLIC and MAPPING packages of the GILDAS<sup>1</sup> software (July 2018 version). We note that the images were corrected for primary beam attenuation. The single-dish data were reduced with the GILDAS-CLASS package.

## 2.3. Missing flux and data merging

To estimate the portion of flux that is missed by the interferometer (due to spatial filtering), we compared the NOEMA and IRAM-30 m data. In this context, the NOEMA data were convolved with a Gaussian beam similar to that of the 30 m data (i.e.,  $\sim 26''$  at 96 GHz) and then smoothed to the same spectral resolution as that of the 30 m observations. By comparing the peak intensities in the direction of the methanol emission peak, we estimate that more than 80% of the methanol emission is resolved out.

To recover the missing flux, we merged the 30 m with the NOEMA data through a routine in the GILDAS-MAPPING package. The resulting data cubes have a velocity resolution of  $0.12 \text{ km s}^{-1}$ . The rms of the resulting spectral data cubes varies from 4 to  $15 \text{ mJy beam}^{-1}$ . The synthesized beam of the combined data cube is  $2.8'' \times 2.4''$  at a position angle of  $29^\circ$ , with a pixel size of  $0.53'' \times 0.53''$ .

## 3. Spatial distribution

### 3.1. Continuum emission

We present in Fig. 1 the continuum maps at 3.1 mm (97 GHz) and 3.6 mm (82 GHz) along with the location of the MMS-1, MMS-2 and MMS-3 sources. Surprisingly, MMS-3 is not detected in our maps, although it is barely detected at the  $3\sigma$  level at 0.87 and 1.2 mm using ALMA observations (Tokuda et al. 2014, 2016) and is detected at the  $5\sigma$  level, as shown in Fig. 2, with ALMA Cycle 3 observations<sup>2</sup> that were carried out at 0.87 mm (project code: ADS/JAO.ALMA#2015.1.00340.S, PI: K. Tokuda. For further details on the data reduction, see Tokuda et al. 2017, 2018).

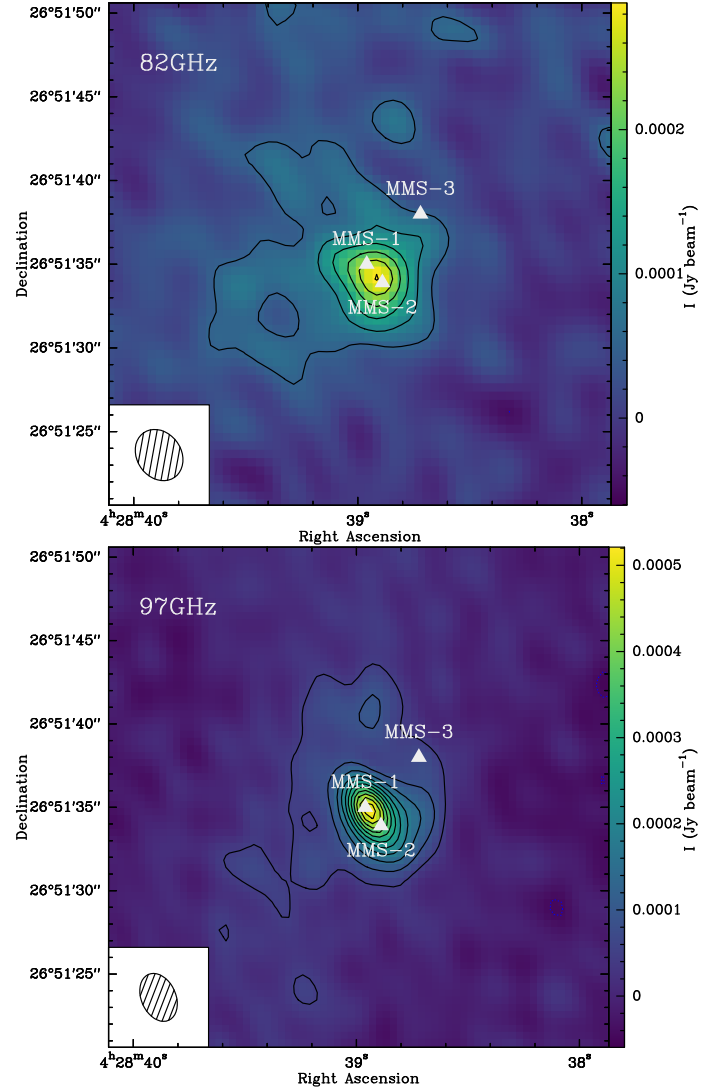
The current resolution of our NOEMA observations does not allow us to distinguish between the positions of sources MMS-1 and MMS-2 (see Fig. 1). The total (MMS-1 + MMS-2) measured flux density per synthesized beam,  $S_\nu$ , is about  $0.03 \text{ mJy beam}^{-1}$  at 82 GHz and  $0.05 \text{ mJy beam}^{-1}$  at 97 GHz. Finally, it is interesting to note that MMS-2 is also detected at 230 GHz in the CALYPSO IRAM-PdBI survey by Maury et al. (2019), but not in the CALYPSO pilot program performed at the same frequency (Maury et al. 2010). We infer that the pilot data along with their calibration and reduction were preliminary.

### 3.2. Methanol channel emission maps

In this subsection, we only present the resulting line emission obtained through the combined NOEMA and IRAM-30 m data for the two following detected methanol transitions:  $E_2 2_{1,2}-1_{1,1}$  at 96.739362 GHz and  $A^+ 2_{0,2}-1_{0,1}$  at 96.741375 GHz. We note

<sup>1</sup> <https://www.iram.fr/IRAMFR/GILDAS/>

<sup>2</sup> The ALMA Cycle 3 continuum data were kindly given to us by K. Tokuda.

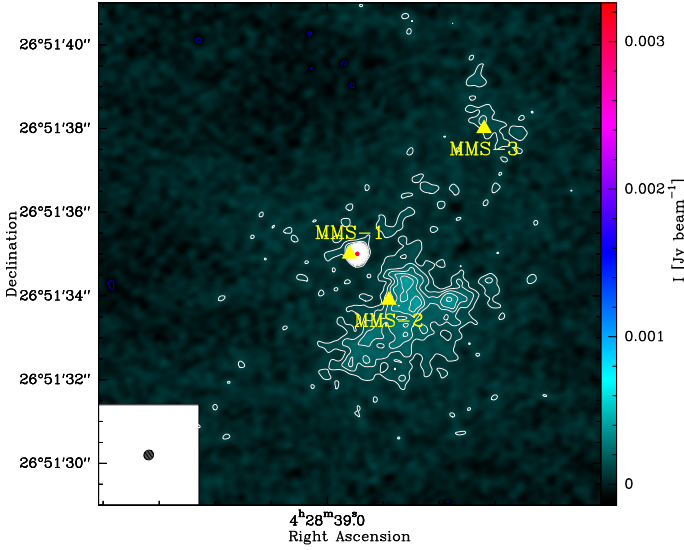


**Fig. 1.** 82 GHz (top panel: 3.65 mm) and 97 GHz (bottom panel: 3.1 mm) continuum emission as observed with NOEMA toward L1521F. The first contour and the level step are at  $3\sigma$  (where  $1\sigma = 1.6 \times 10^{-5}$  and  $1.5 \times 10^{-5} \text{ Jy beam}^{-1}$  at 82 and 97 GHz, respectively). The white triangles indicate the positions of the MMS-1, MMS-2, and MMS-3 sources (see Sect. 1). Synthesized beams are shown in the bottom left corner of the panels.

that the higher energy level  $E_1 2_{0,2}-1_{0,1}$  line at 96.744550 GHz line ( $E_{\text{up}} = 20.1 \text{ K}$ ) is marginally detected (with an rms level of  $3.6 \text{ mJy beam}^{-1}$  or  $0.08 \text{ mK}$ ).

Figure 3 shows the channel emission maps for the  $A^+$ - and  $E_2$ -CH<sub>3</sub>OH lines. The respective emission presents an arc-like structure. A similar filamentary or arc-like structure has previously been observed in this source at the same scale by Tokuda et al. (2014) for HCO<sup>+</sup> ( $J = 3-2$ ). In addition, at about  $5-6 \text{ km s}^{-1}$ , the  $^{12}\text{CO}$  ( $J = 3-2$ ) emission traces an arc-like filamentary structure around L1521F (Tokuda et al. 2016) that is similar to that seen in HCO<sup>+</sup> and methanol (see Fig. 6 from Tokuda et al. 2016, and Fig. 3). It is interesting to note that a CH<sub>3</sub>OH ring-like distribution is also seen in other prestellar cores, such as TUKH122 (which is on the verge of star formation, see Ohashi et al. 2018). In this context, Tafalla et al. (2004) suggest that the ring-like morphology for CH<sub>3</sub>OH in prestellar cores is due to depletion of C-bearing species close to the dust emission peak.





**Fig. 2.** 0.87 mm continuum map observed with ALMA (ADS/JAO.ALMA#2015.1.00340.S, PI: K. Tokuda). The first contour is at  $3\sigma$  and the level step at  $2\sigma$  (where  $1\sigma = 3.3 \times 10^{-5}$  Jy beam $^{-1}$ ). The synthesized beam is shown in the bottom left corner. The positions of sources MMS-1, MMS-2, MMS-3 reported by Tokuda et al. (2014) are indicated as yellow triangles.

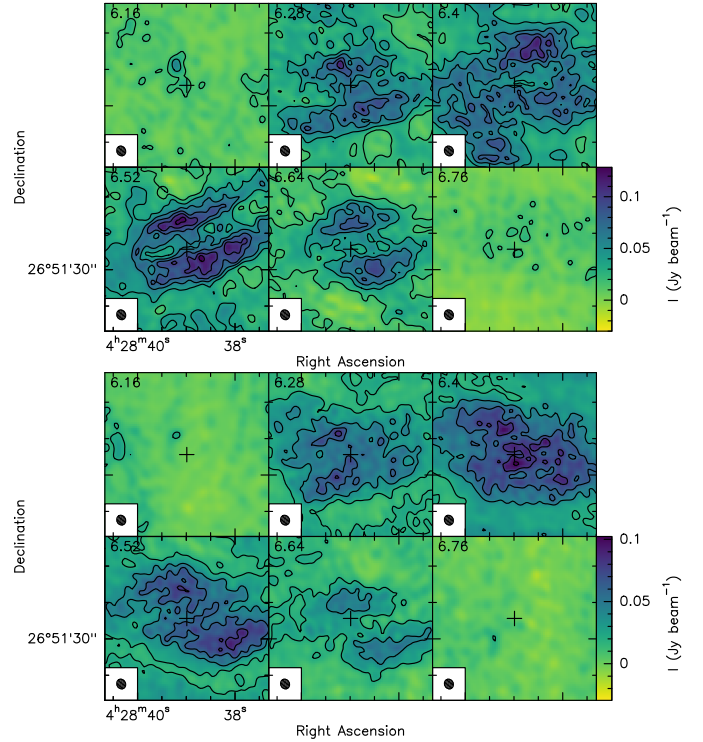
Interestingly enough, Punanova et al. (2018) has observed a centrally peaked emission fragment for CH<sub>3</sub>OH around the center of the L1544 prestellar core and inferred that the methanol emission could arise from an accretion shock. Such structures are likely the result of dynamical gas interaction such as fragmentation (see Tokuda et al. 2014). In this context, we note that similar structures have been reproduced by hydrodynamical simulations with and without magnetic field (Matsumoto et al. 2015, 2017). Turbulence, injected by protostellar feedback, may indeed play a crucial role during fragmentation, different from what can be found in massive disks (Larson 1987; Machida et al. 2008).

### 3.3. Distribution of the methanol emission

We present in Fig. 4 the integrated intensities of the A<sup>+</sup>, E<sub>1</sub>, and E<sub>2</sub> methanol lines (left panel) observed with NOEMA alone. The CH<sub>3</sub>OH emission is compact and peaks toward the northeast position of L1521F at coordinates  $\alpha_{2000} = 04^h28^m39.164^s$ ,  $\delta_{2000} = 26^\circ51'41.49''$ ). Surprisingly, this position is not associated with any of the three MMS sources. The source size is about 7'', corresponding to a 950 au size at a distance of 136 pc ( $\sim 5 \times 10^{-3}$  pc).

In Fig. 4 (right panels) we present the combined 30 m-NOEMA images of the CH<sub>3</sub>OH lines (see Sect. 2.3). The maps show that methanol is indeed extended and distributed in a ring-like structure around the *Spitzer* continuum source MMS-1, with  $V_{LSR} = 6.4$  km s $^{-1}$ . The emission is still brightest at the location of the CH<sub>3</sub>OH peak that appears in the NOEMA-only images. This methanol peak (hereafter called methanol blob) resembles the methanol emission peak found toward the L1544 prestellar core, which is also located in the Taurus molecular cloud (hence at the same distance as L1521F; see Bizzocchi et al. 2014). Incidentally, we note that in the Taurus molecular cloud-1 the CH<sub>3</sub>OH peak is also shifted from the denser part (Soma et al. 2015).

Interestingly enough, Tokuda et al. (2018) have recently focused on the large-scale morphology and kinematics of the



**Fig. 3.** Velocity-channel maps of the A<sup>+</sup> (top) and E<sub>2</sub>-CH<sub>3</sub>OH (bottom) emission. The lowest contour starts at 20 mJy beam $^{-1}$ , with a step of 20 mJy beam $^{-1}$ .

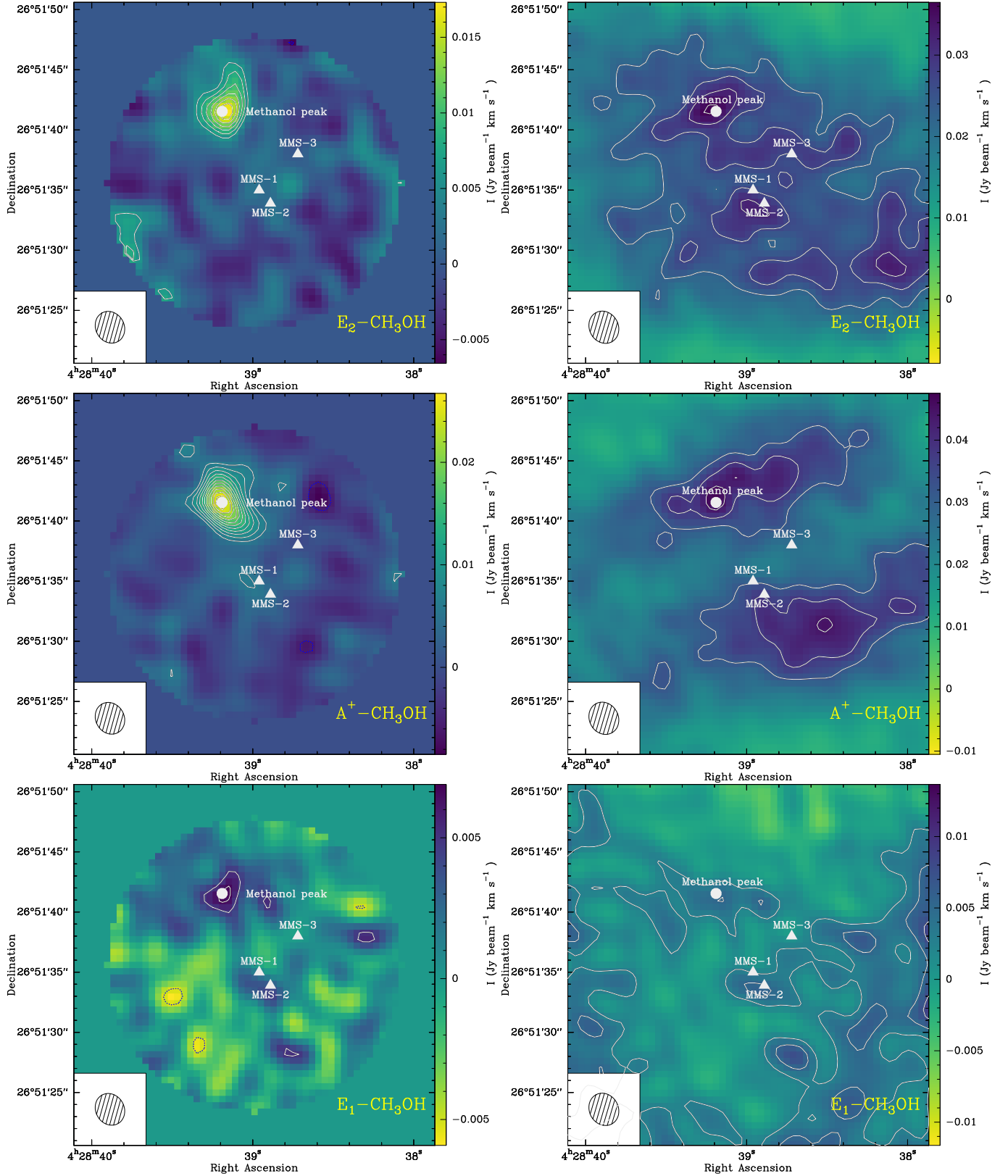
molecular gas around the protostar to understand the dynamical nature of the system, using <sup>12</sup>CO (3–2), <sup>12</sup>CO (2–1), and C<sup>18</sup>O (2–1). From their analysis, MMS-2 is located southwest of the protostar in a warm filament with a 60 K kinetic temperature at a velocity range of 4.45–5.30 km s $^{-1}$ . The emission probed with NOEMA seems to be located at the intersection between three thin, cold (10–30 K), and dense filaments ( $n \sim 10^6$  cm $^{-3}$ ). We cannot exclude that the observed methanol blob is part of the filamentary structure seen in CO, and it might be the result of accreting material.

### 3.4. Distribution of the CS and SO emission

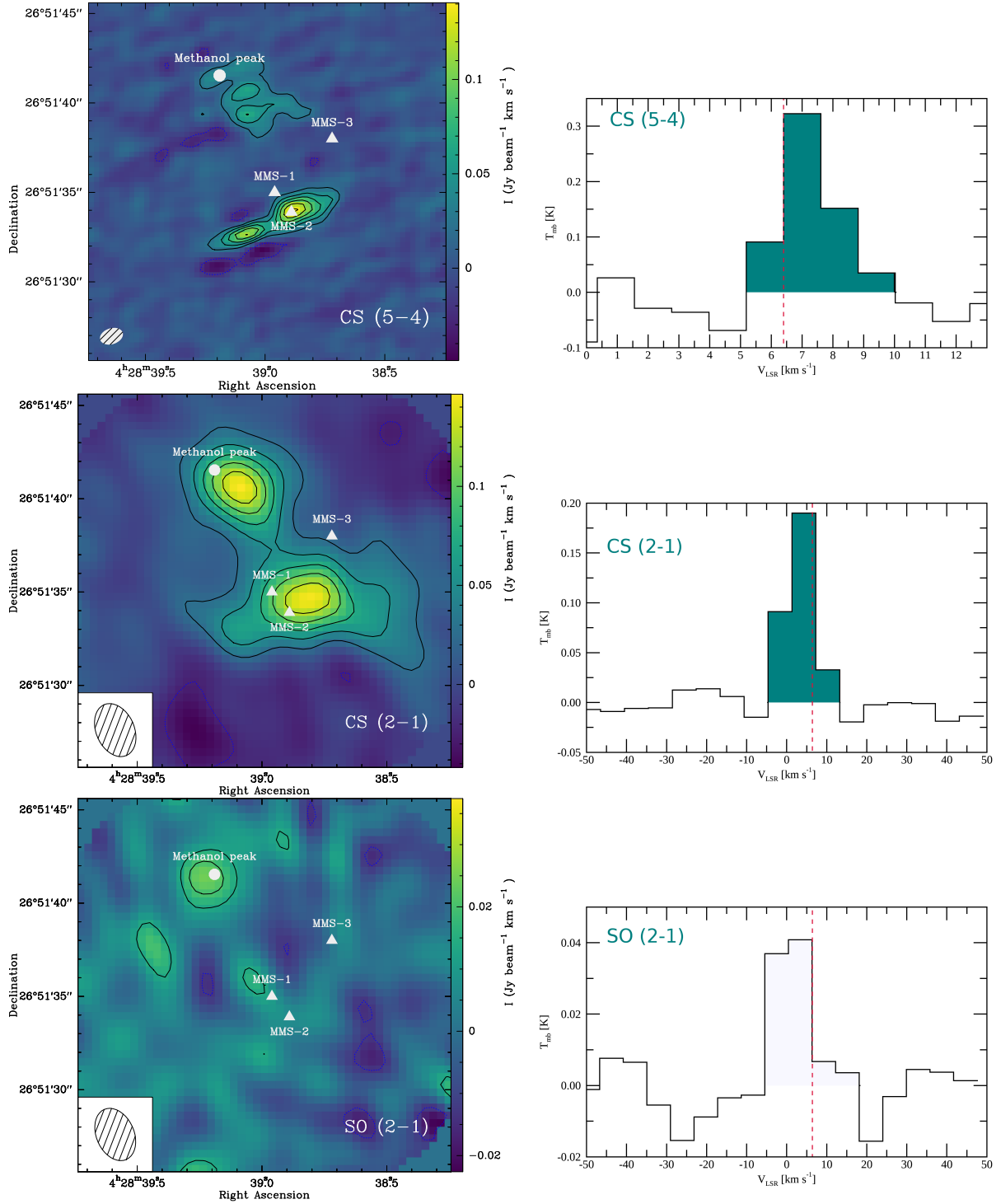
CS (2–1) at 97.980 GHz and SO (2<sub>3</sub>–1<sub>2</sub>) at 99.299 GHz were observed within the WideX bandwidth and the integrated-intensity maps are shown in Fig. 5. CS is clearly detected at the position of the methanol blob and at the position of the VeLLO, while SO is only tentatively detected toward the methanol fragment. The CS emission is compact ( $\sim 5''$ ). In Table 1 we list the spectroscopic and observed parameters of the CS (2–1) line integrated over 5''. CS (5–4) has also been mapped and detected with ALMA toward L1521F (Tokuda et al. 2014). We again retrieved these data from the ALMA data archive (project number 2012.1.00239.S; Early Cycle 0 data performed with a 12 m array) and extracted the spectrum within the same circular beam of 5'' toward the position of the methanol blob (see Table 1 for the measured parameters for this line). Both detections indicate that the gas at this position is very dense (see the following section for the radiative transfer).

### 3.5. Unidentified transition

Finally, we report the detection of an unidentified line (U-line) at the rest frequency of 97 200.541 MHz. Three clumps can



**Fig. 4.** Methanol integrated-emission maps (corrected from the primary beam attenuation). *Top left:*  $E_2$ -CH<sub>3</sub>OH integrated intensity emission map over the line profile. The first contour is at  $3\sigma$  and the level step at  $1\sigma$  (where  $1\sigma = 1.9 \text{ mJy beam}^{-1} \text{ km s}^{-1}$ ). *Top right:* CH<sub>3</sub>OH- $E$  moment-zero maps from the combined IRAM-30 m and NOEMA data. The first contour is at  $5\sigma$  and the level step at  $1\sigma$  (where  $1\sigma = 4.5 \text{ mJy beam}^{-1} \text{ km s}^{-1}$ ). *Middle left:*  $A^+$ -CH<sub>3</sub>OH integrated-intensity emission map over the line profile. The first contour is at  $3\sigma$  and the level step at  $1\sigma$  (where  $1\sigma = 1.9 \text{ mJy beam}^{-1} \text{ km s}^{-1}$ ). *Middle right:*  $A^+$ -CH<sub>3</sub>OH moment-zero maps from the combined IRAM-30 m and NOEMA data. The first contour is at  $5\sigma$  and the level step at  $1\sigma$  (where  $1\sigma = 6.3 \text{ mJy beam}^{-1} \text{ km s}^{-1}$ ). *Bottom left:*  $E_1$ -CH<sub>3</sub>OH integrated-intensity emission map over the line profile. The first contour is at  $3\sigma$  and the level step at  $1\sigma$  (where  $1\sigma = 1.6 \text{ mJy beam}^{-1} \text{ km s}^{-1}$ ). *Bottom right:*  $E_1$ -CH<sub>3</sub>OH moment-zero maps from the combined IRAM-30 m and NOEMA data. The first contour is at  $1\sigma$  ( $2.6 \text{ mJy beam}^{-1} \text{ km s}^{-1}$ ).



**Fig. 5.** *Top left panel:* CS (5–4) integrated-intensity emission map as seen with ALMA (project number 2012.1.00239.S, see Sect. 3.4). The contour levels are at  $3\sigma$  (where  $1\sigma = 7.5$  mJy beam<sup>-1</sup> km s<sup>-1</sup>). *Middle left panel:* CS (2–1) integrated-intensity emission map as seen with NOEMA over the line profile (corrected for primary beam attenuation). The contour levels are at  $3\sigma$  (where  $1\sigma = 8.3$  mJy beam<sup>-1</sup> km s<sup>-1</sup>). *Bottom left panel:* SO (2–1) integrated-intensity emission map as seen with NOEMA over the line profile. The first contour is at  $2\sigma$  and the level step at  $1\sigma$  (where  $1\sigma = 6.6$  mJy beam<sup>-1</sup> km s<sup>-1</sup>). For each map, the synthesized beam is shown in the bottom left corner. Positions of sources MMS-1, MMS-2, and MMS-3 along with that of the methanol peak (or blob) are indicated. *Right panels, from top to bottom:* spectra of the CS (5–4), CS (2–1), and SO (2–1) spectra taken in direction of the methanol blob. Dashed red lines indicate a  $V_{LSR} = 6.4$  km s<sup>-1</sup>.

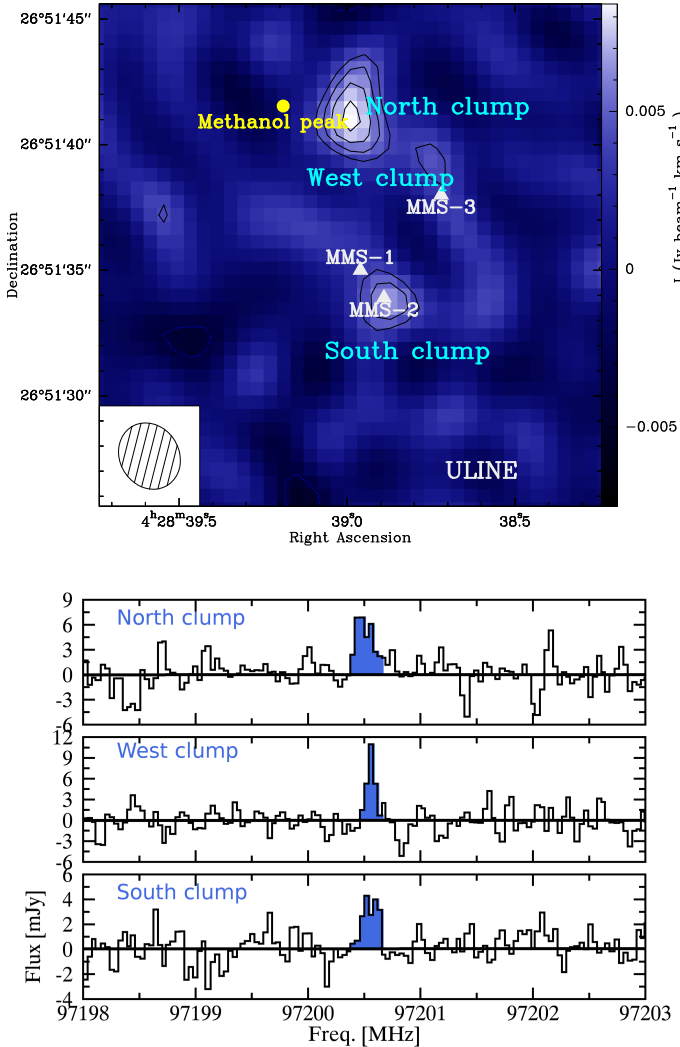
be resolved in the map seen in Fig. 6: north (corresponding to the methanol peak), west (corresponding to MMS-3), and south (corresponding to MMS-1 and MMS-2). We verified the line identification using the CASSIS software and the CDMS and JPL databases. The line parameters are described in Table 2.

The only possible candidate corresponds to vinyl alcohol, a-H<sub>2</sub>C=CHOH ( $5_{3,3} - 4_{3,2}$ , tag = 44 507 in the CDMS database) with  $\nu = 97\,200.6$  MHz,  $E_{up} = 36.95$  K  $A_{ij} = 9.3 \times 10^{-7}$  s<sup>-1</sup>, which has been detected in SgrB2 only by Turner & Apponi (2001). In that instance, the a-H<sub>2</sub>C=CHOH ( $2_{2,1} - 3_{1,2}$ ) transition

**Table 1.** Line parameters measured for CS with NOEMA and ALMA toward the position of the methanol peak and integrated over 5'' using a Gaussian line-fitting procedure from the CASSIS software.

Species	Transition	Frequency (MHz)	$v$ (km s <sup>-1</sup> )	$\Delta v$ (km s <sup>-1</sup> )	Intensity K	Integrated flux (K km s <sup>-1</sup> )
CS	2–1	97980.9533	5.77 (0.24)	10.98 (0.52)	0.12 (0.01)	1.32 (0.17)
CS	5–4	244935.5565	5.92 (0.20)	2.86 (0.58)	0.11 (0.02)	0.31 (0.12)

**Notes.** CASSIS (Centre d'Analyse Scientifique de Spectres Instrumentaux et Synthétiques): <http://cassis.irap.omp.eu>. The numbers in parentheses refer to the  $1\sigma$  level uncertainty derived from the Gaussian fit.



**Fig. 6.** *Top panel:* U-line integrated-intensity emission map over the line profile. The first contour is at  $3\sigma$  and the level step at  $1\sigma$  (where  $1\sigma = 1.3$  mJy beam<sup>-1</sup> km s<sup>-1</sup>). The synthesized beam is shown in the bottom left corner. Positions of sources MMS-1, MMS-2, and MMS-3 along with that of the methanol peak are indicated along with that of the three main U-line emission peaks. *Bottom panel:* spectra of the U-line at 97200.541 MHz taken in the direction of the main emission peaks.

at 96745.9 MHz (with  $E_{up} = 12.99$  K and  $A_{ij} = 1.1 \times 10^{-6}$  s<sup>-1</sup>) should also appear in the other NOEMA sub-bands, which is not the case. Therefore, this species can then be dismissed on the basis that only one transition has been detected. We also verified that this U-line is not a remnant from the other sideband. This line does not appear in the IRAM-30 m data, probably due to

heavy beam dilution. We also explored the L1544 NOEMA data (Punanova et al. 2018) for this transition, which is not detected.

## 4. Molecular column densities and abundances

### 4.1. Carbon monosulfide

CS is an excellent probe of the molecular gas density. We used both 2–1 and 5–4 transitions (see Sect. 3) to constrain the kinetic temperature and density of the gas in the methanol fragment. The spectral resolution for both transitions is unfortunately very different, leading to a much larger line width for the 2–1 transition ( $\sim 11$  km s<sup>-1</sup>, as observed with WideX) compared to the line width of  $\sim 3$  km s<sup>-1</sup> for the 5–4 transition (measured with ALMA), as shown in Fig. 5. This discrepancy in the full width at half-maximum (FWHM) is due to the use of data that were observed with different spectral resolutions: the ALMA data were performed with a spectral resolution of 1.4 km s<sup>-1</sup>, while the NOEMA data (WideX correlator) were observed with a spectral resolution of 6.4 km s<sup>-1</sup>. We used the collisional rates of CS with H<sub>2</sub> calculated by Lique et al. (2006) for temperatures in the range from 10 to 300 K. We first performed a local thermal equilibrium (LTE) analysis using both CASSIS and MADCUBA (Martín et al. 2019), assuming that the LTE approximation holds as the derived densities are high. We then assumed a line width of about 3 km s<sup>-1</sup> and used non-LTE radiative transfer modeling using Radex (van der Tak et al. 2007) within CASSIS. The CS data are consistent with a kinetic temperature range between 10 and 20 K for a density range  $[5 \times 10^5 - 3 \times 10^6]$  cm<sup>-3</sup> and a column density of  $[5.5 - 6.5] \times 10^{12}$  cm<sup>-2</sup>. These high densities and cold kinetic temperatures suggest that a cold and dense condensation formed within the L1521F star-forming system.

In a second step we used the line intensity of the central CS (2–1) channel, which has a width of 3 km s<sup>-1</sup>. This is similar to the CS (5–4) measured line width. In this way, we avoid comparing the full integrated intensity of the CS (2–1) lines with the narrower CS (5–4) line. The minimum  $T_k$  is about 15 K with  $N(\text{CS}) = 1.8 \times 10^{12}$  cm<sup>-2</sup> and  $n(\text{H}_2) > 10^7$  cm<sup>-3</sup>. Using the same calculation but the line intensity of the CS (2–1) line of the two adjacent channels and the CS (5–4)  $3\sigma$  rms noise level in the spectra, we obtain that  $N(\text{CS}) < 8.5 \times 10^{11}$  cm<sup>-2</sup> and  $n(\text{H}_2) < 4 \times 10^5$  cm<sup>-3</sup> for the gas at velocities with no CS (5–4) detections. This implies that the density of the methanol blob is a factor 25 higher (at least) than the gas density in the surrounding environment.

### 4.2. Methanol

Figure 7 shows the averaged spectrum in a 7'' beam around the methanol peak or blob using the NOEMA narrow correlator unit (in Kelvin and Jansky). The spectrum is centered at the frequency of the strongest line ( $A^+ - \text{CH}_3\text{OH}$ , with  $E_{up} = 6.97$  K)



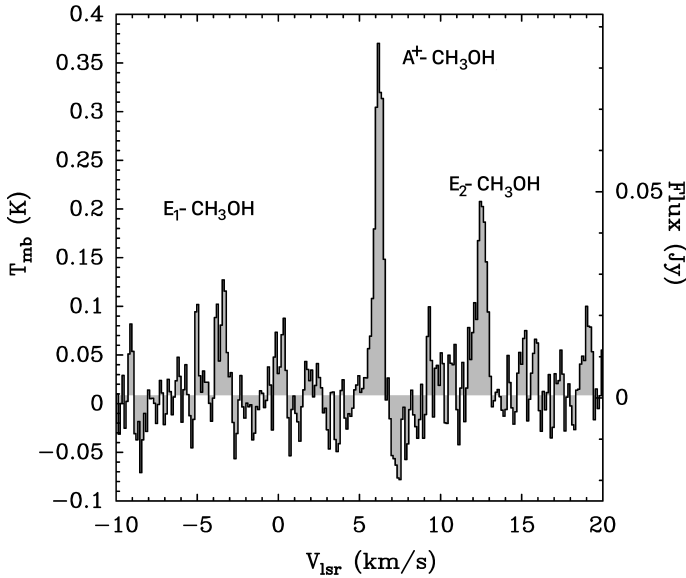
**Table 2.** Location of the U-line and results from the Gaussian line-fitting using the CASSIS software.

Clump	RA (J2000)	Dec (J2000)	<i>FWHM</i> (km s <sup>-1</sup> )	Intensity (mK)	rms (mK)
North clump	04:28:38.99	26:51:41.13	$0.59 \pm 0.12$	230	62
West clump	04:28:38.743	26:51:39.28	$0.25 \pm 0.04$	360	64
South clump	04:28:38.879	26:51:33.84	$0.48 \pm 0.10$	210	56

**Table 3.** Line parameters for CH<sub>3</sub>OH as observed with the narrowband correlator toward the blob integrated over 7".

Frequency (MHz)	$\nu$ (km s <sup>-1</sup> )	$\Delta\nu$ (km s <sup>-1</sup> )	Intensity K	Integrated flux (K km s <sup>-1</sup> )
96739.36	6.31 (0.04)	0.63 (0.04)	0.20 (0.01)	0.13 (0.02)
96741.37	6.21 (0.02)	0.56 (0.04)	0.35 (0.02)	0.20 (0.02)
96744.55	6.38 (0.06)	0.6 (fixed)	$\leq 0.10$	$\leq 0.06$ (3 $\sigma$ )

**Notes.** The numbers in parentheses refer to the  $1\sigma$  level uncertainty derived from the Gaussian fit.

**Fig. 7.** Spectrum of the three methanol transitions in a 7" circular beam around the methanol peak or blob using the NOEMA narrow correlator unit. The spectrum is centered at the frequency of the strongest component (A<sup>+</sup>–CH<sub>3</sub>OH) at 96.74137 GHz.

at 96.74137 GHz. The E<sub>2</sub>–CH<sub>3</sub>OH transition at 96.73936 GHz ( $E_{\text{up}} = 4.64$  K) is also clearly detected at 96.73936 GHz, but the 96.74454 GHz E<sub>1</sub> transition ( $E_{\text{up}} = 12.19$  K) is marginally detected at the  $3\sigma$  level with a peak intensity of about 0.1 K. The lines can be fit with a Gaussian fit with a  $V_{\text{LSR}}$  of 6.2 km s<sup>-1</sup> and an FWHM of  $0.60 \pm 0.05$  km s<sup>-1</sup>. We present the results from the Gaussian line fitting carried out within the CASSIS software in Table 3.

Based on the densities derived from the CS transitions ( $n(\text{H}_2) > 10^7$  cm<sup>-3</sup>), the methanol lines are most likely in LTE. We performed a simple LTE radiative transfer modeling on the two CH<sub>3</sub>OH detected transitions as well as on the upper limit, and found an excitation temperature range of  $(10 \pm 2)$  K and a column density of  $(1.3 \pm 0.2) \times 10^{13}$  cm<sup>-2</sup>. The resulting excitation temperature is compatible with the kinetic temperature obtained from the non-LTE analysis of the CS transitions.

## 5. Discussion

### 5.1. Comparison with previous observations

The high central density and infall asymmetry seen in the HCO<sup>+</sup>(3–2) line observed toward L1521F indicate an object in the earliest stages of gravitational collapse (Onishi et al. 1999). Detection of a 100 au scale dust-continuum source with 1.3 mm PdBI observations (Maury et al. 2010) supports the claim that the protostar has already formed at the center of L1521F. Single-dish (Caltech Submillimeter Observatory) studies in the CO (7–6 and 6–5) emission detected warm ( $\sim 30$ –70 K) and extended ( $\sim 2400$  au) gas, suggesting that this emission may be originating in shocked gas at the interface between the outflow and dense core (Shinnaga et al. 2009).

Tokuda et al. (2014) carried out ALMA Cycle 0 observations toward the object at an angular resolution of  $\sim 1''$  using the 12 m array, revealing that the spatial and velocity distributions are very complex. They detected a few starless high-density clumps ( $\sim 10^7$  cm<sup>-3</sup>), within a region of several hundred au around the *Spitzer* source, a very compact bipolar outflow centered at the protostar source with a dynamical time of a few hundred years with an indication of interaction of surrounding gas and a well-defined long arc-like structure whose size is  $\sim 2000$  au. More recent ALMA Cycle 1 observations have been performed by Tokuda et al. (2016) with a sub-arcsecond resolution, leading to the detection of three intensity peaks at 0.87 mm (MMS-1:  $\alpha_{2000} = 04^{\text{h}}28^{\text{m}}38.96^{\text{s}}$ ,  $\delta_{2000} = 26^{\circ}51'35''$ , MMS-2:  $\alpha_{2000} = 04^{\text{h}}28^{\text{m}}38.89^{\text{s}}$ ,  $\delta_{2000} = 26^{\circ}51'33.9''$ , and MMS-3:  $\alpha_{2000} = 04^{\text{h}}28^{\text{m}}38.72^{\text{s}}$ ,  $\delta_{2000} = 26^{\circ}51'38''$ ). MMS-1 corresponds to the *Spitzer* source L1521F-IRS. Their CO (3–2) and HCO<sup>+</sup> (3–2) observations reveal a complex structure that links all three 0.87 mm peaks as well as the L1521F-NE source detected by Takahashi et al. (2013) with the SMA interferometer. The CO blueshifted and redshifted components observed by the SMA are distributed symmetrically and seem to result from multiple outflows from a binary system, one associated with L1521F-IRS and another associated with this new source, L1521F-NE. However, no driving source has been detected in either millimeter continuum emission with PdBI/SMA or infrared emission with *Spitzer*. Tokuda et al. (2016) were skeptical of the outflows identified by Takahashi et al. (2013) because the direction of the outflow is inconsistent with the *Spitzer* reflection



nebula and suggests instead that it is a relatively high-density gas structure surrounding L1521F. The molecular line observation showed several cores with arc-like structures, possibly due to the dynamical gas interaction. Similar arc-like structures have been reproduced by hydrodynamical simulations with and without a magnetic field (Matsumoto et al. 2015, 2017). The complex structure indicates that in this source turbulence, probably injected by the protostellar feedback, may play an essential role in undergoing fragmentation in the central part of the cloud core. The mechanism is different from the classic scenarios of fragmentation in massive disks (Larson 1987; Boss 2002; Machida et al. 2008). All the above single-dish and interferometric observations demonstrated that significant temperature variations are identified within the core, justifying the need for a high spatial resolution of the central regions.

### 5.2. CH<sub>3</sub>OH ring-like structure in L1521F

The CH<sub>3</sub>OH emission peak detected in the L1521F cluster appears at a distance about 1000 au from the center of the core. In analogy to the L1544 prestellar core, the CH<sub>3</sub>OH peak found with NOEMA toward L1521F resembles the CH<sub>3</sub>OH peak reported at ~4000 au toward the northeast of L1544 (Bizzocchi et al. 2014). In addition, as shown in Fig. 4, L1521F also shows a ring-like structure in CH<sub>3</sub>OH around the MMS-1 source, which is also similar to that observed toward the L1544 prestellar core (Bizzocchi et al. 2014). The CH<sub>3</sub>OH peak and ring in L1544 coincides with the region in the core where CO starts to freeze out and deuterium fractionation starts to be enhanced (Caselli et al. 1999, 2002).

This ring-like morphology of the CH<sub>3</sub>OH emission in L1521F and L1544 has also been reported in other prestellar cores (see also Tafalla et al. 2004), and it is likely the result of several factors: (i) the depletion of C-bearing species as the density increases with decreasing radii within the core; (ii) nonthermal desorption processes such as chemical reactive desorption; (iii) the photo-destruction of CH<sub>3</sub>OH at visual extinctions  $A_V \leq 5$  mag in the outskirts of the core (Vasyunin et al. 2017); and (iv) sputtering from a gentle shock (see Sect. 5.3). As a consequence, it is expected that molecular complexity is high in the external layers of prestellar cores, as confirmed observationally in L1544 (Vastel et al. 2014, 2016, 2018, 2019; Jiménez-Serra et al. 2016; Quénard et al. 2017). At the location of the methanol peak, Jiménez-Serra et al. (2016) found oxygen-bearing complex organic molecules such as CH<sub>3</sub>CHO, HCOOCH<sub>3</sub>, and CH<sub>3</sub>OCH<sub>3</sub>, as well as methoxy (CH<sub>3</sub>O), all related to the release of methanol in the gas phase (Balucani et al. 2015; Soma et al. 2015; Bertin et al. 2016; Vasyunin et al. 2017).

### 5.3. On the nature of the CH<sub>3</sub>OH blob in L1521F

The NOEMA-only maps of CH<sub>3</sub>OH obtained toward L1521F resolve out extended methanol emission and reveal that the methanol peak or blob is very compact (950 au). The physical properties derived for this blob are  $T_k \sim (10 \pm 2)$  K and  $n(\text{H}_2) > 10^7 \text{ cm}^{-3}$ . When we compare these values to those of the L1544 prototypical prestellar core (i.e.,  $n(\text{H}_2) \sim 10^6 \text{ cm}^{-3}$  and  $T_k \sim 7$  K for a 500 au radius; see Crapsi et al. 2007 and Fig. 2 in Vastel et al. 2018), we find that they are quite similar. However, while the methanol peak in L1544 is located at ~4000 au from the core center, the methanol blob in L1521F is found at roughly ~1000 au. At this distance, methanol is clearly depleted in L1544 (Bizzocchi et al. 2014; Vastel et al. 2014; Jiménez-Serra

et al. 2014; Punanova et al. 2018) due to the high density and low temperatures found at this distance in the core.

Crapsi et al. (2004) used the 1.2 mm continuum data of L1521F from the IRAM-30 m to estimate the density distribution under the assumption of spherical symmetry. They followed the same technique as adopted by Tafalla et al. (2002) and best-fit their data with a model of the form

$$n(r) = \frac{10^6}{1 + \left(\frac{r}{20''}\right)^2}. \quad (1)$$

With this density profile, the expected density at the location of the methanol blob is about  $9 \times 10^5 \text{ cm}^{-3}$ , lower than the density estimated in Sect. 4.2 (higher than  $10^7 \text{ cm}^{-3}$ ) from the excitation analysis of the CS (2–1) and (5–4) lines. The resulting derived density appears to be higher than that expected from the  $n(\text{H}_2)$  gas density distribution; this suggests that the CH<sub>3</sub>OH blob might have undergone a compression event of some sort.

In this context, recent NOEMA observations of L1544 within the SOLIS large program focused on the small-scale morphology of the methanol peak emission (Punanova et al. 2018). The kinetic temperature and H<sub>2</sub> gas column density measured for the methanol peak in L1544 from the NOEMA data are 10 K and  $(2.3 \pm 0.3) \times 10^{22} \text{ cm}^{-2}$ . Punanova et al. (2018) concluded that this local methanol enhancement could be an indication of gentle accretion of material onto the core or an interaction of two filaments that produce a slow shock. The methanol peak emission in L1544 is much more extended (more than 20'') than the peak emission detected in the L1521F region (~5'') and appears closer to the center of the core. It is interesting to note that no thermal continuum emission is detected toward the methanol blob. Therefore the fragmentation scenario (see Sect. 5.1) might also explain our observations.

Finally, as briefly mentioned in the previous section, the very presence of methanol in the gas phase in such a cold (~10 K) environment is itself a strong message on its origin. Specifically, because methanol is believed to be a grain-surface product (e.g., Watanabe & Kouchi 2002; Rimola et al. 2014) and the temperature is too low for thermal desorption to play a role, some other mechanism is at work. A first mechanism might be photodesorption from UV photons that penetrate up to the methanol blob, but laboratory experiments suggest that the iced methanol would be injected into the gas phase only as fragments (such as CH<sub>3</sub>O) and not as whole molecules (Bertin et al. 2016). A second often evoked mechanism is the so-called chemical desorption, which is the idea that the energy of the grain-surface reaction is partially transmitted to the product, in this case methanol, to desorb it. While this mechanism could be valid for some species (see, e.g., Oba et al. 2018), it does not seem efficient for methanol according to laboratory experiments (Minissale et al. 2016; Chuang et al. 2018). However, if the composition of the icy mantles includes a higher concentration of CO, methanol could be efficiently chemically desorbed (see Vasyunin et al. 2017). A last possibility is represented by the sputtering caused by a gentle shock (e.g., Flower & Pineau des Forets 1995). This last hypothesis seems to be the most probable for the following reasons: (i) the location of the methanol blob, which does not coincide with any continuum emission peak; (ii) the high density ( $\geq 10^7 \text{ cm}^{-3}$ ), which is higher than the surrounding density by more than a factor 25; (iii) the relatively small extent (~5''), which indicates a very localized phenomenon. If the methanol blob is due to such a gentle shock, then the presence of methanol in the gas-phase would also be easier to explain.

If the methanol blob is due to a gentle shock, then the latter is extremely recent because methanol would very quickly freeze out back onto the grain mantles; this would take only a few hundred years. This could also explain why SO is not detected in our observations. If SO, as is commonly assumed, is formed in the gas phase by oxidation of sulfur that is released from the grain mantles in the form of S or other hydrogenated, organo, or metallic S-bearing species (e.g., [Laas & Caselli 2019](#)), SO would need a few thousand years to form (depending on the gas temperature history: e.g., [Wakelam et al. 2004](#); [Vidal & Wakelam 2018](#)).

In summary, under the hypothesis that the methanol blob is recent, at most a few hundred years, a shock would likely explain the presence of methanol and the absence of SO in the gas phase. The origin of this shock could be a channel of infalling gas toward the center of L1521F. Alternatively, we cannot exclude the hypothesis of the formation of a cold and dense methanol fragment as a result of gas dynamics.

## 6. Conclusions

The original goal of the SOLIS IRAM-NOEMA large program to detect several crucial organic molecules in a sample of solar-like star-forming regions in different evolutionary stages and environments is not achieved for the L1521F very low luminosity object. Instead, we revealed for the first time the presence of a methanol blob emission in the northeastern part of the region, which is located at about  $\sim 1000$  au away from the L1521F source. Our study suggests that at the intersection of a filamentary system (studied previously with ALMA in  $\text{HCO}^+$  and CO) we observe either the formation (i) of a shock-induced cold dense blob or (ii) that of a cold dense fragment. Further observations are needed to distinguish between the two scenarios.

Finally, these observations took place before the implementation of the wideband high-performance correlator PolyFiX that achieved a much higher sensitivity and a much larger bandwidth. A follow-up study at the IRAM 30m will be presented in a forthcoming paper, with deuterated species as well as COMs.

**Acknowledgements.** We thank our referee, Dr. Kazuki Tokuda, (i) for his fruitful comments that have improved the quality of our paper and (ii) for sharing his continuum emission map. This work is supported by the French National Research Agency in the framework of the Investissements d'Avenir program (ANR-15-IDEX-02), through the funding of the "Origin of Life" project of the Univ. Grenoble-Alpes. C.F., C.V. and C.C. acknowledge the funding from the European Research Council (ERC) under the European Unions Horizon 2020 research and innovation programme, for the Project The Dawn of Organic Chemistry (DOC), grant agreement No 741002. I.J.-S. has received partial support from the Spanish FEDER (project number ESP2017-86582-C4-1-R), and State Research Agency (AEI) through project number MDM-2017-0737 Unidad de Excelencia María de Maeztu–Centro de Astrobiología (INTA-CSIC). A.P. acknowledges the financial support of the Russian Science Foundation project 18-12-00351. A.C.-T acknowledges support from MINECO project AYA2016-79006-P.

## References

André, P., Arzoumanian, D., Könyves, V., Shimajiri, Y., & Palmeirim, P. 2019, *A&A*, **629**, L4  
 André, P., Men'shchikov, A., Bontemps, S., et al. 2010, *A&A*, **518**, L102  
 Balucani, N., Ceccarelli, C., & Taquet, V. 2015, *MNRAS*, **449**, L16  
 Bertin, M., Romanzin, C., Doronin, M., et al. 2016, *ApJ*, **817**, L12  
 Bizzocchi, L., Caselli, P., Spezzano, S., & Leonardo, E. 2014, *A&A*, **569**, A27  
 Boss, A. P. 2002, *ApJ*, **568**, 743  
 Bourke, T. L., Myers, P. C., Evans, Neal J., I., et al. 2006, *ApJ*, **649**, L37  
 Caselli, P., Walmsley, C. M., Tafalla, M., Dore, L., & Myers, P. C. 1999, *ApJ*, **523**, L165  
 Caselli, P., Walmsley, C. M., Zucconi, A., et al. 2002, *ApJ*, **565**, 344  
 Ceccarelli, C., Caselli, P., Fontani, F., et al. 2017, *ApJ*, **850**, 176  
 Chen, X., Arce, H. G., Zhang, Q., et al. 2010, *ApJ*, **715**, 1344

Chen, X., Arce, H. G., Dunham, M. M., et al. 2012, *ApJ*, **751**, 89  
 Chuang, K. J., Fedoseev, G., Qasim, D., et al. 2018, *A&A*, **617**, A87  
 Codella, C., Welser, R., Henkel, C., Benson, P. J., & Myers, P. C. 1997, *A&A*, **324**, 203  
 Commerçon, B., Launhardt, R., Dullemond, C., & Henning, T. 2012, *A&A*, **545**, A98  
 Crapsi, A., Caselli, P., Walmsley, C. M., et al. 2004, *A&A*, **420**, 957  
 Crapsi, A., Caselli, P., Walmsley, C. M., et al. 2005, *ApJ*, **619**, 379  
 Crapsi, A., Caselli, P., Walmsley, C. M., & Tafalla, M. 2007, *A&A*, **470**, 221  
 Dunham, M. M., Evans, Neal J., I., Bourke, T. L., et al. 2006, *ApJ*, **651**, 945  
 Dunham, M. M., Crapsi, A., Evans, Neal J., I., et al. 2008, *ApJS*, **179**, 249  
 Dunham, M. M., Chen, X., Arce, H. G., et al. 2011, *ApJ*, **742**, 1  
 Enoch, M. L., Lee, J.-E., Harvey, P., Dunham, M. M., & Schnee, S. 2010, *ApJ*, **722**, L33  
 Flower, D. R., & Pineau des Forets, G. 1995, *MNRAS*, **275**, 1049  
 Hosokawa, T., Omukai, K., Yoshida, N., & Yorke, H. W. 2011, *Science*, **334**, 1250  
 Jiménez-Serra, I., testi, L., Caselli, P., & Viti, S. 2014, *ApJ*, **787**, L33  
 Jiménez-Serra, I., Vasyunin, A. I., Caselli, P., et al. 2016, *ApJ*, **830**, L6  
 Laas, J. C., & Caselli, P. 2019, *A&A*, **624**, A108  
 Larson, R. B. 1969, *MNRAS*, **145**, 271  
 Larson, R. B. 1987, *Am. Sci.*, **75**, 376  
 Lee, C. W., Bourke, T. L., Myers, P. C., et al. 2009, *ApJ*, **693**, 1290  
 Lique, F., Spielteufel, A., & Cernicharo, J. 2006, *A&A*, **451**, 1125  
 Machida, M. N., Inutsuka, S.-i., & Matsumoto, T. 2008, *ApJ*, **676**, 1088  
 Maheswar, G., Lee, C. W., & Dib, S. 2011, *A&A*, **536**, A99  
 Martín, S., Martín-Pintado, J., Blanco-Sánchez, C., et al. 2019, *A&A*, **631**, A159  
 Masunaga, H., & Inutsuka, S.-i. 2000, *ApJ*, **531**, 350  
 Masunaga, H., Miyama, S. M., & Inutsuka, S.-i. 1998, *ApJ*, **495**, 346  
 Matsumoto, T., Onishi, T., Tokuda, K., & Inutsuka, S. I. 2015, *MNRAS*, **449**, L123  
 Matsumoto, T., Tokuda, K., Onishi, T., et al. 2017, *J. Phys. Conf. Ser.*, **837**, 012009  
 Maury, A. J., André, P., Hennebelle, P., et al. 2010, *A&A*, **512**, A40  
 Maury, A. J., André, P., Testi, L., et al. 2019, *A&A*, **621**, A76  
 Minissale, M., Dulieu, F., Cazaux, S., & Hocuk, S. 2016, *A&A*, **585**, A24  
 Mizuno, A., Onishi, T., Hayashi, M., et al. 1994, *Nature*, **368**, 719  
 Murillo, N. M., & Lai, S.-P. 2013, *ApJ*, **764**, L15  
 Oba, Y., Tomaru, T., Lamberts, T., Kouchi, A., & Watanabe, N. 2018, *Nat. Astron.*, **2**, 228  
 Ohashi, S., Sanhueza, P., Sakai, N., et al. 2018, *ApJ*, **856**, 147  
 Omukai, K. 2007, *PASJ*, **59**, 589  
 Onishi, T., Mizuno, A., Kawamura, A., Ogawa, H., & Fukui, Y. 1996, *ApJ*, **465**, 815  
 Onishi, T., Mizuno, A., Kawamura, A., Ogawa, H., & Fukui, Y. 1998, *ApJ*, **502**, 296  
 Onishi, T., Mizuno, A., Kawamura, A., & Fukui, Y. 1999, in *Star Formation*, ed. T. Nakamoto (Nobeyama: NRO), 153  
 Onishi, T., Mizuno, A., Kawamura, A., Tachihara, K., & Fukui, Y. 2002, *ApJ*, **575**, 950  
 Punanova, A., Caselli, P., Feng, S., et al. 2018, *ApJ*, **855**, 112  
 Quéward, D., Vastel, C., Ceccarelli, C., et al. 2017, *MNRAS*, **470**, 3194  
 Rimola, A., Taquet, V., Ugliengo, P., Balucani, N., & Ceccarelli, C. 2014, *A&A*, **572**, A70  
 Schnee, S., Sadavoy, S., Di Francesco, J., Johnstone, D., & Wei, L. 2012, *ApJ*, **755**, 178  
 Shinnaga, H., Phillips, T. G., Furuya, R. S., & Kitamura, Y. 2009, *ApJ*, **706**, L226  
 Soma, T., Sakai, N., Watanabe, Y., & Yamamoto, S. 2015, *ApJ*, **802**, 74  
 Tafalla, M., Myers, P. C., Caselli, P., Walmsley, C. M., & Comito, C. 2002, *ApJ*, **569**, 815  
 Tafalla, M., Myers, P. C., Caselli, P., & Walmsley, C. M. 2004, *Ap&SS*, **292**, 347  
 Takahashi, S., Ohashi, N., & Bourke, T. L. 2013, *ApJ*, **774**, 20  
 Tokuda, K., Onishi, T., Saigo, K., et al. 2014, *ApJ*, **789**, L4  
 Tokuda, K., Onishi, T., Matsumoto, T., et al. 2016, *ApJ*, **826**, 26  
 Tokuda, K., Onishi, T., Saigo, K., et al. 2017, *ApJ*, **849**, 101  
 Tokuda, K., Onishi, T., Saigo, K., et al. 2018, *ApJ*, **862**, 8  
 Tomida, K., Machida, M. N., Saigo, K., Tomisaka, K., & Matsumoto, T. 2010, *ApJ*, **725**, L239  
 Turner, B. E., & Apponi, A. J. 2001, *ApJ*, **561**, L207  
 van der Tak, F. F. S., Black, J. H., Schöier, F. L., Jansen, D. J., & van Dishoeck, E. F. 2007, *A&A*, **468**, 627  
 Vastel, C., Ceccarelli, C., Lefloch, B., & Bachiller, R. 2014, *ApJ*, **795**, L2  
 Vastel, C., Ceccarelli, C., Lefloch, B., & Bachiller, R. 2016, *A&A*, **591**, L2  
 Vastel, C., Quéward, D., Le Gal, R., et al. 2018, *MNRAS*, **478**, 5514  
 Vastel, C., Loison, J. C., Wakelam, B., & Lefloch, B. 2019, *A&A*, **625**, A91  
 Vasyunin, A. I., Caselli, P., Dulieu, F., & Jiménez-Serra, I. 2017, *ApJ*, **842**, 33

- Vidal, T. H. G. & Wakelam, V. 2018, *MNRAS*, **474**, 5575
- Vorobyov, E. I., Elbakyan, V., Dunham, M. M., & Guedel, M. 2017, *A&A*, **600**, A36
- Wakelam, V., Caselli, P., Ceccarelli, C., Herbst, E., & Castets, A. 2004, *A&A*, **422**, 159
- Watanabe, N., & Kouchi, A. 2002, *ApJ*, **571**, L173
- Young, C. H., Jørgensen, J. K., Shirley, Y. L., et al. 2004, *ApJS*, **154**, 396
- 
- <sup>1</sup> Univ. Grenoble Alpes, CNRS, Institut de Planétologie et d'Astrophysique de Grenoble (IPAG), 38000 Grenoble, France  
e-mail: cecile.favre@univ-grenoble-alpes.fr
- <sup>2</sup> IRAP, Université de Toulouse, CNRS, CNES, UPS, (Toulouse), France
- <sup>3</sup> Centro de Astrobiología (CSIC, INTA), Ctra. de Torrejón a Ajalvir km4, Torrejón de Ardoz, 28850 Madrid, Spain
- <sup>4</sup> School of Physics and Astronomy, Queen Mary University of London, 327 Mile End Road, London, E1 4NS, UK
- <sup>5</sup> Max-Planck-Institut für extraterrestrische Physik (MPE), 85748 Garching, Germany
- <sup>6</sup> Observatorio Astronómico Nacional (OAG-IGN), Alfonso XII 3, 28014 Madrid, Spain
- <sup>7</sup> INAF – Osservatorio Astrofisico di Arcetri, Largo E. Fermi 5, 50125 Florence, Italy
- <sup>8</sup> Department of Physics and Astronomy, University College London, Gower Street, WC1E 6BT, London, UK
- <sup>9</sup> Leiden Observatory, Leiden University, PO Box 9513, 2300 RA Leiden, The Netherlands
- <sup>10</sup> Department of Physics, The University of Tokyo, Bunkyo-ku, 113-0033 Tokyo, Japan
- <sup>11</sup> Ural Federal University, 620002, 19 Mira street, Yekaterinburg, Russia
- <sup>12</sup> RIKEN Cluster for Pioneering Research, 2-1, Hirosawa, Wako-shi, 351-0198 Saitama, Japan
- <sup>13</sup> Institut de Radioastronomie Millimétrique, 300 rue de la Piscine, Domaine Universitaire de Grenoble, 38406 Saint-Martin d'Hères, France
- <sup>14</sup> Dipartimento di Chimica, Biologia e Biotecnologie, Via Elce di Sotto 8, 06123 Perugia, Italy
- <sup>15</sup> Departament de Química, Universitat Autònoma de Barcelona, 08193 Bellaterra, Catalonia, Spain
- <sup>16</sup> LERMA, Université de Cergy-Pontoise, Observatoire de Paris, PSL Research University, CNRS, Sorbonne Université, UPMC, Univ. Paris 06, 95000 Cergy Pontoise, France
- <sup>17</sup> University of AL-Muthanna, College of Science, Physics Department, AL-Muthanna, Iraq
- <sup>18</sup> Univ. Rennes, CNRS, IPR (Institut de Physique de Rennes) – UMR 6251, 35000 Rennes, France
- <sup>19</sup> ESO, Karl Schwarzschild Str. 2, 85478 Garching bei München, Germany
- <sup>20</sup> Aix-Marseille Université, PIIM UMR-CNRS 7345, 13397 Marseille, France
- <sup>21</sup> Università degli Studi di Torino, Dipartimento Chimica Via Pietro Giuria 7, 10125 Torino, Italy
- <sup>22</sup> Engineering Research Institute Ventpils International Radio Astronomy Centre of Ventpils University of Applied Sciences, Inženieru 101, Ventpils 3601, Latvia