

THE HERSCHEL REFERENCE SURVEY

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ABSTRACT

The Herschel Reference Survey is a guaranteed time Herschel key project and will be a benchmark study of dust in the nearby universe. The survey will complement a number of other Herschel key projects including large cosmological surveys that trace dust in the distant universe. We will use Herschel to produce images of a statistically-complete sample of 323 galaxies at 250, 350 and 500 μm . The sample is volume-limited, containing sources with distances between 15 and 25 Mpc and flux limits in the K-band to minimize the selection effects associated with dust and with young high-mass stars and to introduce a selection in stellar mass. The sample spans the whole range of morphological types (ellipticals to late-type spirals) and environments (from the field to the centre of the Virgo Cluster) and as such will be useful for other purposes than our own. We plan to use the survey to investigate (i) the dust content of galaxies as a function of Hubble type, stellar mass and environment, (ii) the connection between the dust content and composition and the other phases of the interstellar medium and (iii) the origin and evolution of dust in galaxies. In this paper, we describe the goals of the survey, the details of the sample and some of the auxiliary observing programs that we have started to collect complementary data. We also use the available multi-frequency data to carry out an analysis of the statistical properties of the sample.

Subject headings: Galaxies: general, ISM; Infrared: galaxies; Submillimeter; Catalogs; Surveys

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1. INTRODUCTION

Understanding the processes that govern the formation and evolution of galaxies is one of the major challenges of modern astronomy. Ideally this work can be done by analyzing and comparing the physical, structural and kinematical properties of different objects at various epochs to model predictions. How primordial galaxies transform their huge gas reservoir into stars and become the objects that we now observe in the local universe is a key question. A wide coverage of the electromagnetic spectrum to probe the atomic (21 cm line) and molecular (generally through the 2.6 mm CO line) gas components, the different stellar populations (UV to near-IR spectral range), the dust (mid- and far-IR) and the electrons in magnetic fields (radio continuum) is necessary to address this important question.

The importance of dust resides in the fact that it is formed from the aggregation of the metals produced in the latest phases of stellar evolution, and contains approximately half the metals in the interstellar medium (Whittet 1992). Injected into the interstellar medium (ISM) by stellar winds and supernovae explosions, dust acts as a catalyst in the process of transformation of the atomic to molecular hydrogen necessary to feed star formation (Hollenbach & Salpeter 1971; Duley & Williams 1986). Dust also contributes to the shielding of the UV radiation field, preventing the dissociation of molecular clouds, and thus playing a major role in the energetic equilibrium of the ISM (Hollenbach & Tielens 1997). Dust contributes to the cooling and heating of the ISM in photodissociation regions through photoelectric heating. Furthermore, by absorbing the stellar light, dust modifies our view of the different stellar components (e.g. Buat & Xu 1996): because dust obscuration is so important in star-forming regions, the emission from dust is one of the most powerful tracers of the star formation activity in galaxies (Kennicutt 1998; Hirashita et al. 2003). Dust emission can therefore be used to study the relationship between the gas surface density and the star formation activity, generally known as the Schmidt law.

The importance of dust in the study of the formation and evolution of galaxies became evident after the IRAS space mission which provided us with a whole sky coverage in four infrared bands. Despite its low sensitivity and poor angular resolution, IRAS detected tens thousands of extragalactic sources (Soifer et al. 1987). The IRAS survey showed that the stellar light in some galaxies is so heavily obscured that the only way to determine their star formation activity is through the dust emission itself. Among these highly obscured galaxies, the Ultraluminous Infrared Galaxies (ULIRGs) are the most actively star-forming objects in the local universe. With their improved sensitivity, spectral and angular resolution, other space missions such as ISO and Spitzer have brought a better understanding of the chemical and physical properties of interstellar dust in a wide range of different galactic and extragalactic sources. IRAS, ISO, Spitzer and the recently launched AKARI satellite, however, are sensitive to dust emitting in the mid- ($\sim 5 \mu\text{m}$) to far- ($\leq 200 \mu\text{m}$) infrared

domain. The 5-70 μm spectral range corresponds to the emission of the hot dust generally associated to star formation, while at longer wavelengths, up to 1 mm, the contribution of cold dust becomes more and more important (Draine et al. 2007). The emission of the ISM in the range $3 \mu\text{m} \lesssim \lambda \lesssim 15 \mu\text{m}$ is generally dominated by Polycyclic Aromatic Hydrocarbons (PAHs). Between 10 μm and $\lesssim 70 \mu\text{m}$ the dust emission is usually due to very small grains, while that at longer wavelengths ($70 \mu\text{m} \lesssim \lambda \lesssim 1000 \mu\text{m}$) is probably produced by big grains of graphite and silicate in thermal equilibrium with the UV and optical photons of the interstellar radiation field (Désert et al. 1991; Dwek et al. 1997; Zubko et al. 2004; Draine & Li 2007).

There is, however, strong empirical evidence to suggest that much of the dust in normal galaxies has been missed by previous space missions because it is too cold to radiate in the mid- and far-infrared. Devereux & Young (1990), for example, showed that when the dust masses of galaxies were estimated from IRAS data, the gas-to-dust ratios were $\simeq 10$ times higher than the Galactic value, implying that 90% of the dust in galaxies is too cold to radiate significantly in the IRAS bands. ISO 200 μm images of nearby galaxies revealed the presence of cold (~ 20 K) dust in the external parts of galaxies (Alton et al. 1998). Even Spitzer, with its longer wavelength coverage, is only sensitive to dust that is warmer than $\simeq 15$ K (Bendo et al. 2003; Draine & Li 2007; Draine et al. 2007). Simple estimates show that the temperature of a dust grain in thermal equilibrium in the average interstellar radiation field is ~ 20 K, which would produce a modified black body spectrum with a peak at $\sim 200 \mu\text{m}$ with a flux rapidly decreasing at longer wavelengths (Boselli et al. 2003; Dale et al. 2005, 2007; Draine & Li 2007; Draine et al. 2007). It is important to remember, however, that most of the dust in a galaxy is likely to be at a much lower temperature than 20 K because it is self-shielded, and so unaffected from the interstellar radiation field. For example, in molecular clouds the temperature of the dust is usually well below the canonical 20 K value once the average dust extinction (A_V) is greater than one, unless it is close to a newly formed star (Ward-Thompson et al. 2002).

The submillimetre waveband (200-1000 μm) is crucial for detecting this missing cold dust component and thus for making accurate estimates of the total dust mass. Mass estimates made from far-infrared measurements are highly uncertain because of the strong dependence of flux on dust temperature which is difficult to determine using data on the Wien side of the black body peak. At long wavelengths (Rayleigh-Jeans domain) the emissivity depends only on the first power of the dust temperature, making it possible to use a submillimetre flux to make an accurate estimate of the mass of dust (Eales et al. 1989; Galliano et al. 2003; 2005).

Submillimetre continuum observations of galaxies have been made previously, in particular with the ground-based submillimetre cameras, such as SCUBA on the James Clerk Maxwell Telescope. These have confirmed that galaxies do indeed contain a large amount of cold dust (Dunne & Eales 2001; Galliano et al. 2003). Until

now, the largest submillimetre survey has been the SCUBA Local Universe and Galaxy Survey (SLUGS), a survey of $\simeq 200$ galaxies in two samples, one selected from the IRAS survey and one selected at optical wavelengths. The survey produced the first estimates of the submillimetre luminosity and dust-mass (the space density of galaxies as a function of dust mass) functions (Dunne et al. 2000; Vlahakis et al. 2005), and also indicated that some ellipticals are surprisingly strong submillimetre sources. Its limitation, and the limitation of all ground-based submillimetre observations, is that of sensitivity: SLUGS struggled to detect galaxies not already detected by IRAS.

The solution is to go into space and the Herschel space telescope is finally making this possible. In particular the relatively large field of view, high sensitivity, and coverage of a waveband (250-500 μm) in which galaxies are intrinsically bright, make SPIRE on Herschel the ideal instrument for a study of all extragalactic sources (Griffin et al. 2006, 2007; Wasket et al. 2007). SPIRE has the sensitivity to map large areas of the sky down to the confusion limit in quite modest integration times, providing submillimetre data for large samples of galaxies at different redshifts, and is thus well adapted for pointed observations for all kind of nearby extragalactic sources.

The SPIRE consortium has defined a number of coordinated guaranteed time programs to get the best possible benefit of the unique Herschel facilities for the study of galaxy evolution. These include deep cosmological surveys, pointed observations of local galaxies and surveys of representative regions in the Milky Way. To better characterize the dust properties in the local universe the SPIRE extragalactic group has selected a volume-limited sample of 323 galaxies to be observed in guaranteed time in the three SPIRE bands at 250, 350 and 500 μm . The importance of the local universe resides in the fact that it represents the endpoint of galaxy evolution, providing important boundary conditions to models and simulations. Furthermore, at ≤ 30 Mpc the angular resolution of SPIRE (a couple of kpc) allow us to resolve the different galaxy components such as nuclei, bulges, discs and spiral arms. Moreover, dwarf galaxies, by far the most common (yet still very poorly understood) galaxies in the universe, can only be observed locally.

A volume-limited sample was chosen as a way to limit distance dependent selection biases. To limit any possible contamination due to the cosmic variance, the volume should be much larger than the typical dimension of large-scale structures (~ 30 Mpc). At the same time it should be representative of the whole galaxy population inhabiting the local universe. A near-infrared K-band selection was adopted for two major reasons: (i) unlike optical surveys, which have strong selection effects due to the presence of dust, near-IR data are only marginally affected by dust extinction; (ii) whereas the optical flux is highly dependent on the number of relatively young stars and thus sensitive to recent episodes of star formation, the near-IR luminosity is a good measure of the overall stellar mass (e.g. Gavazzi

et al. 1996), which recent studies suggest as the most important parameter in characterizing the statistical and evolutionary properties of galaxies. Indeed galaxy properties that appear to be tightly correlated with the galaxy's stellar mass include the following: physical properties, such as the star formation activity, the gas content and the metallicity (Boselli et al. 2001; 2002; Zaritsky et al. 1994; Gavazzi et al. 2004; Tremonti et al. 2004); structural properties, such as the concentration index and the galaxy's light profile (Boselli et al. 1997; Gavazzi et al. 2000; Scodreggio et al. 2002); the amount and distribution of dark matter, as indicated by the Tully-Fisher relation and the shape of the rotation curve for spirals (Tully & Fisher 1977; Catinella et al. 2006) and the fundamental plane for ellipticals (Dressler et al. 1987; Djorgovski & Davis 1987); the stellar population, shown through the colour-magnitude relations for early-type (Visvanathan & Sandage 1977, Bower et al. 1992) and late-type galaxies (Tully et al. 1982, Gavazzi et al. 1996) and the galaxy spectral energy distributions (Gavazzi et al. 2002, Kauffmann et al. 2003). These correlations all indicate a down-sizing effect (e.g. Gavazzi et al. 1996; Cowie et al. 1996; Heavens et al. 2004), in which galaxies with high stellar masses formed most of their stars at a much earlier cosmic epoch than those with low stellar masses. This underline the dominant role of mass, rather than morphology, in shaping galaxies.

The final sample includes galaxies with a large range in luminosity (mass) and includes all Hubble types. Because of its definition, the selected sample also includes galaxies belonging to different density regions such as the core of the Virgo cluster, groups and pairs and isolated objects. Given its completeness, the Herschel Reference Sample (HRS) will be a suitable reference for any statistical study. Combined with KINGFISH (the Herschel extension of SINGS, Kennicutt et al. 2003) and VNGS (see below), two samples optimized for the study of the different phases of the ISM in individual galaxies, these samples will provide the community with a unique dataset for studying the ISM properties of galaxies in the local universe.

The paper is organized as follows. Section 2 describes the scientific goals of the survey. Section 3 gives a description of the selection criteria used to define the sample. Section 4 gives details of the SPIRE observations we will carry out and Section 5 gives an overview of the multi-frequency data that is available for the sample. Section 6 gives a brief description of the data processing and products. The multifrequency statistical properties of the sample will be described in Section 7.

2. SCIENTIFIC OBJECTIVES

The overall goal of this Herschel survey is to improve our knowledge of the cold dust properties of the most common extragalactic sources populating the local universe. Combined with multifrequency ancillary data covering the whole electromagnetic spectrum (see Sect. 5), the new Herschel data will provide us with a unique dataset with which to:

- 1) Trace, for the first time, the variation of

the properties of the cold dust component (dust mass and temperature, dust to gas ratio...) of normal galaxies along the Hubble sequence, and as a function of luminosity. Given the large dispersion in galaxy properties (Morton & Haynes 1994), it is important the sample to be large enough to contain representatives of all galaxy types and include both early- and late-type galaxies spanning the largest possible range in luminosity. Where galaxies are resolved, we can analyze the distribution of the cold dust in the different galaxy components, e.g. the nuclei, bulges, spiral arms and diffuse discs. This will provide important constraints on dust formation and evolution in galaxies (Galliano et al. 2008).

2) Study the role of dust in the physics of the ISM. As discussed in the introduction, dust plays a major role in the energetic equilibrium of the ISM. It acts as catalyst for the formation of the molecular hydrogen (Hollenbach & Salpeter 1971; Duley & Williams 1986) and shields the molecular gas component preventing dissociation from the diffuse interstellar radiation field (Hollenbach & Tielens 1997). To understand the nature of the ISM we thus need to know how the cold dust properties (temperature, composition, geometrical distribution...) relate to other physical properties such as metallicity and interstellar radiation field (Boselli et al. 2004; Galliano et al. 2008). By combining dust surface brightnesses with metallicity dependent gas to dust ratios and HI column densities, sub-mm measurements can be used to determine H₂ column densities (Guélin et al. 1993; 1995; Neinger et al. 1996; Boselli et al. 2002). SPIRE data will therefore provide us with an independent measure of the molecular hydrogen component, and can be use for an accurate calibration of the still highly uncertain CO to H₂ conversion factor.

3) Study the relationship between dust emission and star formation. Dust participates in the cooling of the gas through the shielding of the interstellar radiation field, in particular of the UV light, and thus plays a major role in the process of star formation (Draine 1978; Dwek 1986; Hollenbach & Tielens 1997). The energy absorbed by dust is then radiated in the infrared domain. For the same reason dust emission has often been used as a tracer of star formation. We still do not know, however, what is the relationship between the cold dust emission and star formation. The study of resolved galaxies will allow us to analyze the relationship between the infrared surface brightness and the dust temperature (Chañal et al. 2007) and to trace the connections between the star formation and the dust and gas properties at galactic scales, extending the recent results of Spitzer to all phases of the ISM (Gordon et al. 2004, Calzetti et al. 2005, 2007; Perez-Gonzalez et al. 2006; Kennicutt et al. 2007, Prescott et al. 2007; Thilker et al. 2007).

4) Study the dust extinction properties in galaxies. Astronomers have tackled with the problem of dust opacity in galaxies for over 40 years (e.g. Holmberg 1958; Disney et al. 1989; Calzetti 2001), but have reached no consensus. The key problem is that optical catalogs contain huge selection effects because of

the existence of dust. We will be able to address this issue in two ways. First, the submillimetre images will, for the first time, allow us to determine the distribution of the dust column density in a very large number of galaxies. Second, we will be able to determine the global dust opacity in each galaxy by using the energy balance between the absorbed stellar light and the dust emitted radiation (Buat & Xu 1996, Witt & Gordon 2000, Buat et al. 2002, Cortese et al. 2006; 2008a). This requires an accurate determination of the UV to near-IR (stellar light) and mid-IR to sub-mm (dust emission) spectral energy distribution (Boselli et al. 2003). By comparing the dust attenuation properties of different classes of objects this analysis will allow us to define standard recipes for correcting UV and optical data, a useful tool for the interpretation of all modern surveys.

5) Determine whether there is an intergalactic dust cycle. Apart from the obvious possibility that dust is ejected from galaxies by the same methods that gas is ejected, such as through interactions with the surrounding environment (see point 9) and starburst-driven 'superwinds' as indeed observed in M82 (Engelbracht et al. 2006), there is also the possibility that dust is ejected from galaxies by radiation pressure (Davies et al. 1998, Meiksin 2009, Oppenheimer & Davé 2008). The ejection of dust from galaxies might explain the huge reservoir of metals found in the intergalactic medium (Rayan-Weber et al. 2006). There is clear evidence for dust in superwinds (Alton et al. 1999) and there is tantalizing evidence from ISO observations for extended distributions of dust around galaxies (Alton et al. 1998; Davies et al. 1999). Observations so far have been limited by either sensitivity (SCUBA) or resolution (ISO and Spitzer). The Herschel Reference Survey will be able to determine whether dust ejection is important because (i) we will be observing several hundred galaxies, so even if these phenomena are rare our sample should contain some examples, and (ii) our observations will have the sensitivity to detect dust well outside the optical disc of each galaxy.

6) Determine the amount of interstellar dust in ellipticals. Very little is known about dust in ellipticals. Despite the stereotype that ellipticals do not contain any dust, the structures seen in optical images imply that at least 50% of ellipticals contain some dust (Ferrarese et al. 2006). The amount of interstellar dust is too small, however, to be easily detected through its far infrared emission; IRAS, for example, detected only about 15% of ellipticals (Bregman et al. 1998). Although Spitzer has now detected many ellipticals, these studies have mostly been of sources known a priori to contain some dust (Kaneda et al. 2007; Temi et al. 2007; Panuzzo et al. 2007) or molecular gas (Young et al. 2008). The SLUGS did contain a small statistically-complete sample of 11 ellipticals, although some of the six 850- μ m detections may well have been of synchrotron rather than dust emission (Vlahakis et al. 2005). Other ellipticals have been detected at 350 μ m by Leeuw et al. (2008). The HRS contains 65 early-type (E, S0 and S0a) galaxies and our observations will have sufficient sensitivity to detect dust masses down to $\sim 10^4 M_{\odot}$. The Herschel Reference Survey will therefore

provide an unambiguous answer to the question of how much dust is in ellipticals.

7) Determine the origin of dust in ellipticals.

The three possible origins of the dust in ellipticals are that (i) it has been produced in the atmospheres of the old evolved stars that dominate the light of ellipticals today, (ii) its origin is external to the elliptical and is the result of a merger, (iii) it is the relic of dust formed during the active star-forming phase early in the history of the galaxy. The distribution of dust is an important clue to understand its true origin. The presence of the first mechanism is suggested by Spitzer mid-infrared spectra of Virgo ellipticals (Bressan et al. 2006) which show the silicate emission produced by dust in circumstellar envelopes of evolved stars. The origin of this feature (predicted by Bressan et al. 1998) is supported by the finding that the mid-infrared emission has the same profile as optical light (Xilouris et al. 2004) in many early-type galaxies. It is unclear, however, whether this is the predominant mechanism producing the interstellar dust in these galaxies. For example, a submillimetre map of the nearby elliptical Centaurus A, has revealed a dusty disc, implying that the dust (and the galaxy itself) has been formed as the result of a merger (Leeuw et al. 2002). Temi et al. (2007) also found little correlation between the Spitzer 70 and 160 μm emission and optical starlight, which also suggests the dust has an external origin.

We will use the Herschel Reference Survey results to investigate the origin of the dust by, first, investigating the detailed morphology of the dust, in particular how it compares with the stellar distribution, and, second, by investigating whether the mass of dust is correlated with other global properties of the galaxy such as stellar mass. It has often been argued that sputtering by the ubiquitous, hot X-ray emitting gas in early galaxies should destroy any dust formed more than 10^8 years in the past (Tielens et al. 1994). However, the fact that dust is seen visually in 50% of ellipticals (Ferrarese et al. 2006) and the recent discovery of small grains, which are preferentially destroyed by sputtering, in ellipticals (Kaneda et al. 2007) suggests that dust is protected in some way. We will investigate whether sputtering from the hot gas destroys the dust grains and will test the internal origin hypothesis by investigating any possible correlation between X-ray excess and dust mass.

8) Measure the local luminosity and dust-mass distributions.

These distributions are important benchmarks for the deep Herschel surveys. These have already been estimated as part of SLUGS (Dunne et al. 2000; Vlahakis et al. 2005): the much greater sensitivity of the Herschel Reference Survey will push these limits down by a factor of ~ 50 .

9) Understand the role of the environment on the evolution of galaxies.

The well-known phenomenon that spiral galaxies in clusters are HI-deficient and have truncated HI discs (eg. Cayatte et al. 1990) demonstrates that the environment of a galaxy can have a strong effect on its interstellar medium. Indeed there is clear evidence for tidally-stripped dust in interacting

galaxies (Thomas et al. 2002) and indications for ram-pressure stripped dust in cluster objects (Boselli & Gavazzi 2006). At the same time the presence of metals (Sarazin 1986) and possibly of dust (Stickel et al. 2002; Montier & Giard 2005) in the diffuse intra-cluster medium has been shown by X-ray and far-IR observations. The comparison of the submillimetre emission of cluster and isolated galaxies within the HRS will allow us to make a detailed study on the effects of the environment on the dust properties of galaxies, and thus understand whether the hot and dense cluster intergalactic medium can be polluted through the gas stripping process of late-type galaxies (Boselli & Gavazzi 2006).

10) Provide a multi-frequency reference sample suitable for any statistical study.

Our aim is that the HRS will be the first large sample of galaxies with observations of all phases of the ISM, as well as measurements over the entire wavelength range of the spectral energy distribution (SED) for each galaxy. The sample will then serve for many purposes e.g. the UV to radio continuum SEDs could, for example, be used to determine the nature of unresolved sources or as templates for estimating photometric redshifts.

3. THE SAMPLE

The Herschel Reference Sample (HRS) has been selected according to the following criteria:

1) **Volume-limited:** A volume limit was imposed to reduce distance uncertainties due to local peculiar motions and ensure the presence of low-luminosity, dwarf galaxies, not accessible at high redshift. By applying a lower distance (D) limit we exclude the very extended sources, the observing of which would be too time consuming²³. We have selected all galaxies with an optical recessional velocity (vel , taken from NED) in between 1050 km s^{-1} and 1750 km s^{-1} that, for $H_0 = 70 \text{ km s}^{-1} \text{ Mpc}^{-1}$ and in the absence of peculiar motions, corresponds to $15 \leq D \leq 25 \text{ Mpc}$ ²⁴. In the Virgo cluster region ($12\text{h} < \text{R.A.}(2000) < 13\text{h}$; $0^\circ < \text{dec} < 18^\circ$), where peculiar motions are dominant, we have included all galaxies with $vel < 3000 \text{ km s}^{-1}$ and belonging to cluster A, the North (N) and East (E) clouds and the Southern extension (S) (17 Mpc) and Virgo B (23 Mpc), where the subgroup membership has been taken from Gavazzi et al. (1999a). W and M clouds objects, at a distance of 32 Mpc, have been excluded²⁵.

2) **K band selection:** Given the expected low dust

²³ A small sample of very nearby and extended galaxies will be observed in detail as part of another guaranteed time project: Physical Processes in the Interstellar Medium of Very Nearby Galaxies.

²⁴ Given the possible discrepancy between optical and HI recessional velocity measurements, heliocentric velocities given in Table 1 can be outside this range.

²⁵ The sample was selected before distances were available on NED. According to these new NED estimates, the distance of the HRS galaxies outside the Virgo cluster are generally included in the $15 \leq D \leq 25 \text{ Mpc}$ range. In the Virgo region, however, our distances are determined according to subgroup membership criteria, and are thus generally different than those given by NED

content of quiescent galaxies, whose emission would be hardly detectable within reasonable integration times, a more stringent limit has been adopted for early-types than for star forming galaxies. Among those galaxies in the required recessional velocity range, we thus selected all late-type spirals and irregulars (Sa-Sd-Im-BCD) with a 2MASS K band total magnitude $K_{Stot} \leq 12$ and all ellipticals and lenticulars (E, S0, S0a) with $K_{Stot} \leq 8.7$.

3) **High galactic latitude:** To minimize galactic cirrus contamination, galaxies have been selected at high galactic latitude ($b > +55^\circ$) and in low galactic extinction regions ($A_B < 0.2$; Schlegel et al. 1998).

The resulting sample is composed of 323 galaxies located in the sky region between $10h17m < R.A.(2000) < 14h43m$ and $-6^\circ < dec < 60^\circ$ (see Fig. 1), of which 65 are early-type (E, S0 and S0a) and 258 are late-type (Sa-Sd-Im-BCD) (see Fig. 2). Figure 2 shows that the only galaxies which are clearly undersampled are blue compact galaxies and dwarf irregulars, the most numerous galaxies in the nearby universe. They will be the targets of another SPIRE key program²⁶. As selected, the sample spans a large range in environment since it includes the Virgo cluster, many galaxy groups and pairs as well as relatively isolated objects (Fig. 1). Using the Virgo cluster membership criteria defined in Gavazzi et al. (1999a), the HRS includes 82 members of cluster A and B (Fig. 2) The other galaxies are members of nearby clouds such as Leo, Ursa Major and Ursa Major Southern Spur, Crater, Coma I, Canes Venatici Spur and Canes Venatici - Camelopardalis and Virgo-Libra Clouds (Tully 1988). As defined, the sample is thus ideal for environmental studies (Cortese & Hughes 2009; Hughes & Cortese 2009). The whole sample is given in Table 1 with columns arranged as follow:

Column 1: Herschel Reference Sample (HRS) name. This is the position of the galaxy in the sample list when sorted by increasing right ascension.

Column 2: Zwicky name, from the Catalogue of Galaxies and of Cluster of Galaxies (CGCG; Zwicky et al. 1961-1968).

Column 3: Virgo Cluster Catalogue (VCC) name, from Binggeli et al. (1985).

Column 4: Uppsala General Catalog (UGC) name (Nilson 1973).

Column 5: New General Catalogue (NGC) name.

Column 6: Index Catalogue (IC) name.

Columns 7 and 8: J2000 right ascension and declinations, taken from NED.

Column 9: Morphological type, taken from NED, or from our own classification if not available.

Column 10: Total K band magnitude (K_{Stot}), from 2MASS (Skrutskie et al. 2006).

Column 11: Optical isophotal diameter (25 mag arcsec⁻²), from NED.

Column 12: Heliocentric radial velocity (in km s⁻¹), from HI data when available, otherwise from NED.

Column 13: Cluster or cloud membership, from Gavazzi

et al. (1999a) for Virgo and Tully (1988) or Nolthenius (1993) whenever available, or from our own estimate otherwise.

Column 14: Pair/group membership, from Karachentsev et al. (1972) or from NED whenever available, or from our own estimate elsewhere. Pair/group membership has been assigned according to the following criteria: close pairs (CP) are those with a nearby companion at a distance less than 50 kpc and a difference in redshift $< 600 \text{ km s}^{-1}$, as in Gavazzi et al. (1999b), while pairs (P) up to 150 kpc. The number indicates whether the galaxy belongs to a triplet (3), quadruplet (4) and quintuplet (5). Groups outside Virgo and its immediate surroundings have been determined by counting the number of bright galaxies²⁷ within 25 arcmin (which, at a median distance of 20 Mpc, corresponds to ~ 150 kpc) and 600 km s^{-1} . Pairs in the Virgo region are only those cataloged by Karachentsev et al. (1972).

Column 15: Galactic extinction A_B (Schlegel et al. 1998).

Column 16: Alternative names.

²⁶ The ISM in Low Metallicity Environments: Bridging the Gap Between Local Universe and Primordial Galaxies

²⁷ For bright galaxies we intend those included in major catalogs such as NGC, UGC, IC or CGCG.

TABLE 1
CONTINUE

HRS	CGCG	VCC	UGC	NGC	CIC	RA(2000)	dec	type	K_S	$D(25)$	vel	memb	groupAB	Other name
						h m s	° ' "		mag	'	km s^{-1}		mag	
271	100015	-	8014	4758	-	125244.04	155055.9	Im::;HII	10.93	3.00	1240	Virgo Out.		0.11
272	71065	2095	8016	4762	-	125256.05	111350.9	SB(r)0 sp;LINER	7.30	8.70	985	Virgo E Cl.	P	0.09 KPG356B
273	15031	-	8020	4771	-	125321.27	011609.0	SAd? sp;NLGN	9.01	4.00	1119	Virgo Out.		0.09
274	15032	-	8021	4772	-	125329.17	021006.0	SA(s)a;LINER;Sy3	8.36	2.90	1038	Virgo Out.		0.12
275	-	-	-	4775	-	125345.70	-063719.8	SA(s)d	9.22	2.14	1566	Virgo Out.		0.15
276	71068	-	8022	4779	-	125350.86	094235.7	SB(rs)bc;Sbrst	9.87	2.10	2832	Virgo Out.		0.09
277	43060	-	-	4791	-	125443.97	080310.7	cl	11.35	1.20	2529	Virgo Out.		0.14
278	71071	-	8032	-	-	125444.19	131414.2	S	10.39	2.75	1121	Virgo Out.		0.15
279	15037	-	8041	-	-	125512.68	000700.0	SB(s)d	11.98	3.10	1321	Virgo Out.		0.10
280	43066	-	8043	4799	-	125515.53	025347.9	S?	9.89	1.60	2802	Virgo Out.		0.15
281	43068	-	8045	-	-	125523.62	075434.0	IBm:	11.82	0.91	2801	Virgo Out.		0.18
282	43069	-	-	4803	-	125533.67	081425.8	Comp	10.71	0.50	2664	Virgo Out.		0.13
283	43071	-	8054	4808	-	125548.94	041814.7	SA(s)cd::;HII	9.04	2.60	760	Virgo Out.		0.16
284	-	-	-	-	3908	125640.62	-073346.1	SB(s)d?;HII	9.10	1.82	1296	Virgo Out.		0.15
285	15049	-	8078	4845	-	125801.19	013433.0	SA(s)ab sp;HII	7.79	5.20	1097	Virgo Out.		0.09
286	71092	-	8102	4866	-	125927.14	141015.8	SA(r)0+ sp;LINER	7.92	6.00	1986	Virgo Out.		0.12
287	15055	-	8121	4904	-	130058.67	-000138.8	SB(s)cd;Sbrst	9.50	2.40	1174	Virgo Out.		0.11
288	-	-	-	4941	-	130413.14	-053305.8	(R)SAB(r)ab::;Sy2	8.22	3.63	1114	Virgo Out.		0.16
289	-	-	-	4981	-	130848.74	-064639.1	SAB(r)bc;LINER	8.49	2.75	1678	Virgo Out.		0.18
290	189037	-	8271	5014	-	131131.16	361654.9	SA? sp	10.11	1.70	1136	Canes Ven. Spur	3	0.03
291	217031	-	8388	5103	-	132030.08	430502.3	Sab	9.49	1.45	1297	Canes Ven. Spur		0.08
292	218010	-	8439	5145	-	132513.92	431602.2	S?;HII;Sbrst	9.33	2.00	1225	Canes Ven. Spur		0.05
293	16069	-	8443	5147	-	132619.71	020602.7	SB(s)dm	9.73	1.91	1092	Virgo Out.		0.12
294	246017	-	8593	-	902	133601.22	495739.0	Sb	10.42	2.19	1608	Canes Ven. Spur		0.05
295	73054	-	8616	5248	-	133732.07	085306.2	(R)SB(rs)bc;Sy2;HII	7.25	1.79	1152	Virgo-Libra Cl.	P	0.11
296	190041	-	8675	5273	-	134208.34	353915.2	SA(s)0;Sy1.9	8.67	2.75	1064	Canes Ven. Spur		0.04 KPG391A
297	246023	-	8711	5301	-	134624.61	460626.7	SA(s)bc: sp	9.11	4.17	1508	Canes Ven. Spur		0.07
298	218047	-	8725	5303	-	134744.97	381816.4	Pec	10.23	0.91	1419	Canes Ven. Spur	CP	0.06 KPG397A
299	45108	-	8727	5300	-	134816.04	035703.1	SAB(r)c	9.50	3.89	1171	Virgo-Libra Cl.		0.10
300	218058	-	8756	-	-	135035.89	423229.5	Sab	10.34	1.70	1354	Canes Ven. Spur		0.06
301	17088	-	8790	5334	4338	135254.46	-010652.7	SB(rs)c:	9.94	4.17	1380	Virgo-Libra Cl.		0.20
302	45137	-	8821	5348	-	135411.27	051338.8	SBbc: sp	10.87	3.55	1443	Virgo-Libra Cl.	5	0.12
303	295024	-	8843	5372	-	135446.01	583959.4	S?	10.65	0.65	1717	Canes Ven-Came. Cl.		0.04
304	46001	-	8831	5356	-	135458.46	052001.4	SABbc: sp;HII	9.64	3.09	1370	Virgo-Libra Cl.	5	0.11
305	46003	-	8838	5360	958	135538.75	045906.2	I0	11.15	2.19	1171	Virgo-Libra Cl.	5	0.13
306	46007	-	8847	5363	-	135607.21	051517.2	I0?	6.93	4.07	1136	Virgo-Libra Cl.	5	0.12
307	46009	-	8853	5364	-	135612.00	050052.1	SA(rs)bc pec;HII	7.80	6.76	1242	Virgo-Libra Cl.	5	0.12
308	46011	-	8857	-	-	135626.61	042348.0	Sb(f)	11.93	0.91	1091	Virgo-Libra Cl.		0.14
309	272031	-	9036	5486	-	140724.97	550611.1	SA(s)m:	11.95	1.86	1383	Canes Ven-Came. Cl.	3	0.09
310	47010	-	9172	5560	-	142005.42	035928.4	SB(s)b pec	9.98	3.72	1718	Virgo-Libra Cl.	3	0.13 ARP286
311	47012	-	9175	5566	-	142019.95	035600.9	SB(r)ab;LINER	7.39	6.61	1492	Virgo-Libra Cl.	3	0.13 ARP286
312	47020	-	9183	5576	-	142103.68	031615.6	E3	7.83	3.55	1482	Virgo-Libra Cl.	3	0.14
313	47022	-	9187	5577	-	142113.11	032608.8	SA(rs)bc:	9.75	3.39	1490	Virgo-Libra Cl.	3	0.18
314	19012	-	9215	-	-	142327.12	014334.7	SB(s)d	10.54	2.19	1389	Virgo-Libra Cl.		0.14
315	220015	-	9242	-	-	142521.02	393222.5	Sc	11.73	5.01	1440	Canes Ven. Spur		0.05
316	47063	-	9308	5638	-	142940.39	031400.2	E1	8.25	2.69	1845	Virgo-Libra Cl.	4	0.14
317	47066	-	9311	-	1022	143001.85	034622.3	S?	11.70	1.10	1716	Virgo-Libra Cl.		0.15
318	47070	-	9328	5645	-	143039.35	071630.3	SB(s)d	9.69	2.40	1370	Virgo-Libra Cl.		0.12
319	75064	-	9353	5669	-	143243.88	095330.5	SAB(rs)cd	10.35	3.98	1368	Virgo-Libra Cl.		0.12
320	47090	-	9363	5668	-	143324.34	042701.6	SA(s)d	11.71	3.31	1583	Virgo-Libra Cl.	P	0.16
321	47123	-	9427	5692	-	143818.12	032437.2	S?	10.54	0.89	1581	Virgo-Libra Cl.		0.16
322	47127	-	9436	5701	-	143911.06	052148.8	(R)SB(rs)0/a;LINER	8.14	4.27	1505	Virgo-Libra Cl.		0.16
323	48004	-	9483	-	1048	144257.88	045324.5	S	9.55	2.24	1640	Virgo-Libra Cl.	CP	0.16

TABLE 2
THE SPIRE PERFORMANCES

λ	250 μm	350 μm	500 μm
Ang. Res.	18.1"	25.2"	36.9"
Lin. Res.	1.7 kpc	2.4 kpc	3.6 kpc
Sensitivity	12 mJy/beam	8 mJy/beam	12 mJy/beam

Note: $1\text{-}\sigma$ sensitivity (instrumental noise) for a standard map (one repeat = two cross-linked scans) as determined during the science verification phase (on board observations). The linear resolution is at 20 Mpc.

4. SPIRE OBSERVATIONS

With its large field of view ($4' \times 8'$), its sensitivity ($S \sim 7$ mJy, 5σ in 1 hour, point source mode), and angular resolution (~ 30 arcsec; see Table 2), SPIRE on Herschel is the ideal instrument for the proposed survey.

In constructing our observing program, we have used different integration times for early- and late-type galaxies because the former are known to contain less dust.

Integration times for early-types were determined by assuming a lower limit of $\sim 10^4 M_{\odot}$ to their total dust mass as determined from IRAS observations (Bregman et al. 1998) or measured from optical absorption line measurements (Goudfrooij et al. 1994; Van Dokkum & Franx 1995; Ferrari et al. 1999). At 20 Mpc, a galaxy with a dust mass of $\sim 10^4 M_{\odot}$ would have a flux density at 250 μm of 11 mJy. We estimate that with the adopted integration times we should detect an elliptical with $\sim 10^4 M_{\odot}$ of dust at the 3σ level.

Integration times for spirals have been chosen so that we should be able to detect dust outside the optical radius²⁸, for which there is evidence from ISO observations (Alton et al. 1998; Davies et al. 1999). By combining ISOPHOT (Alton et al. 1998) and SCUBA (Vlahakis et al. 2005) observations of extended sources with the spectral energy distribution of normal galaxies of different type (Boselli et al. 2003), and assuming a standard gas to dust ratio of ~ 100 with the extragalactic calibration for the dust-emissivity coefficient (James et al. 2002), we estimate that the dust associated with the extended HI disc would have a flux density of ~ 14 mJy per beam at 250 μm for HI column densities of $\sim 10^{20}$ atoms cm^{-2} . In normal galaxies such gas column densities are generally reached well outside the stellar disc, at ~ 1.5 optical radii (Cayatte et al. 1994; Broelis & Van Woerden 1994). Because of metallicity gradients, however, the gas to dust ratio is expected to increase in the outer discs, as observed in M31 by Cuillandre et al. (2001). We would therefore expect to observe lower far-IR flux densities. With 12 minutes of integration time (3 scan-maps) we will get to an instrumental noise of 6.9 mJy/beam, when the expected emission is 14 mJy per beam; this sensitivity will allow us to detect the dust emission in the outer disc by integrating flux densities along elliptical annuli with a sensitivity of ~ 30 better than that of previous observations (Alton et al. 1998).

All galaxies outside the Virgo cluster will be observed in scan-map mode ($30''/\text{sec}$). For each galaxy, the size of the scan-map has been chosen based on the optical size of the galaxy. For late-type galaxies, the observing mode has been chosen so that our images will cover the galaxy at least out to $1.5 \times D_{25}$ (where D_{25} is the 25 mag arcsec^{-2} isophotal diameter) i.e. out to the approximate HI radius. For spirals we will make three pairs of cross-linked observations, which will reach a 1σ sensitivity in instrumental noise of 6.9, 4.6 and 6.9 mJy/beam at 250, 350 and 500 μm respectively as determined during the science verification phase. These values are comparable to the noise from the confusion of faint unrelated sources (4, 5 and 6 mJy/beam in the three bands). Elliptical galaxies are likely to be weaker and thus need longer integration times, we have compromised by only requiring the image to contain the galaxy out to D_{25} . For ellipticals and S0s we will make eight pairs of cross-linked observations reaching a 1σ sensitivity in instrumental noise of 4.2, 2.8 and 4.2 mJy/beam at 250, 350 and 500 μm respectively, which in the latter two bands is below the noise expected from the confusion of faint sources.

Sixty square degrees centered on the Virgo cluster region will be covered by Herschel in parallel mode (i.e. using both SPIRE and PACS) as part of the Herschel Virgo Cluster Survey HeViCS²⁹. Eighty-three HRS galaxies fall into this region. To avoid duplication, these galaxies will be observed only during the HeViCS survey. HeViCS will survey this region in fast-scan mode ($60''/\text{sec}$) using both SPIRE and PACS. The survey will make eight scans reaching a 1σ of instrumental noise of 4.5, 6.2 and 5.3 mJy/beam at 250, 350 and 500 μm respectively, as determined using the Herschel Observation Planning Tool (HSpot v4.2.0). This is not as deep as our observations of ellipticals and S0s outside Virgo, but the effect of confusion noise means that the decrease in effective sensitivity is really small. On the plus side, for the HRS galaxies in Virgo we will also have PACS images at 110 and 160 μm .

5. COMPLEMENTARY DATA

A number of key aims of our survey (see Section 2) can only be achieved through the combination of Herschel observations with corollary data. UV to near-IR imaging is needed to determine the distribution and content of the different stellar populations, to quantify the intensity of the interstellar radiation field and to study the recent activity of star formation (UV and $\text{H}\alpha$). Optical spectroscopy is crucial to measure stellar and gas metallicities and to determine the presence of any nuclear activity. At the same time, the Balmer emission line can be used to measure the dust extinction within HII regions and are thus essential for quantifying the current level of star formation activity. HI and CO line data are necessary for determining the content and distribution of the gaseous component of the ISM, the principal feeder of star formation in galaxies. High resolution HI imaging can also be used to study the kinematical properties of the target galaxies. Mid-, far-IR and sub-mm data, combined with SPIRE data,

²⁸ Defined as the B band isophotal radius at 25 mag arcsec^{-2}

²⁹ www.arcetri.astro.it/hevics

will be used to study the physical properties of the different dust components (PAHs, very small grains and big grains), and radio continuum data for measuring the thermal and synchrotron emission.

Given its definition, the HRS is easily accessible for ground based and space facilities: the selected galaxies are in fact relatively bright ($m_B < 15$ mag) and extended (~ 2 -3 arcmin). Listed below are the most important references for ancillary data.

5.1. *UV, optical, near-IR and H α imaging*

Of the 323 galaxies in the HRS, 280 have been observed by the Galaxy Evolution Explorer (GALEX) in the two UV bands FUV ($\lambda_{\text{eff}} = 1539\text{\AA}$, $\Delta\lambda = 442\text{\AA}$) and NUV ($\lambda_{\text{eff}} = 2316\text{\AA}$, $\Delta\lambda = 1069\text{\AA}$). Observations have been taken as part of the Nearby Galaxy Survey (NGS; Gil de Paz et al. 2007), the Virgo cluster survey (Boselli et al. 2005), the All Imaging Survey (AIS) or as pointed observations in open time proposals. A GALEX legacy survey has been recently accepted to complete the UV observation of the HRS galaxies at a deepness of the Medium Imaging Survey (1500 sec/field).

SDSS (DR7 release, Adelman-McCarthy et al. 2008) images in the *ugriz* filters are available for 313 objects. All galaxies have been observed in the JHK filters by 2MASS (Jarrett et al. 2003). Deep B,V, H and K band images for all Virgo cluster galaxies and for some Coma I Cloud galaxies are available on the GOLDMine database (<http://goldmine.mib.infn.it/>; Gavazzi et al. 2003). These have been taken with pointed observations during the near infrared and optical extensive surveys of the Virgo cluster carried out by Boselli et al. (1997, 2000; 2003) and Gavazzi et al. (2000).

An H α + [NII] imaging survey of the star forming HRS galaxies found outside the Virgo cluster under way at the 2.1m San Pedro Martir telescope is almost complete. Combined with images taken in the Virgo cluster region (Boselli & Gavazzi 2002, Boselli et al. 2002 and Gavazzi et al. 1998; 2002; 2006) H α + [NII] data are now available for 221 of the 258 late-type objects and 26 of the 65 early-types.

5.2. *Integrated spectroscopy*

A low resolution ($R \sim 1000$), integrated spectroscopic survey in the wavelength range 3500-7200 \AA is under way at the 1.93m OHP telescope. In order to sample the spectral properties of the whole galaxy, and not just those of the nuclear regions (these last provided by the SDSS for 106 galaxies in the DR6), observations have been executed using the drifting technique described in Kennicutt (1992). Exposures are taken while constantly and uniformly drifting the spectrograph slit over the full extent of the galaxy. A resolution $R \sim 1000$ is mandatory for resolving [NII] from H α and measuring the stellar underlying Balmer absorption under H β .

Data for 64 Virgo cluster galaxies in the HRS have already been published in Gavazzi et al. (2004) along with a few other galaxies in Moustakas & Kennicutt (2006), Jansen et al. (2000) and Kennicutt (1992b). We have

data from the literature or from our own observations for 256/258 of the late-types and 33/65 of the early-types, making the whole sample complete at 89 %. The remaining objects will be included in the future observing runs.

5.3. *HI and CO lines*

Single beam HI observations are available for a large variety of sources. Most of the galaxies have HI data in Springob et al. (2005) and Gavazzi et al. (2005), this last reference limited to the Virgo cluster region. Out of the 8 late-type galaxies, 249 have an HI measurements, as do 55/65 of the early-types. All galaxies in the $0^\circ < \text{dec} < 30^\circ$ range (80 %) will be observed during the ALFALFA survey under way at the 305m Arecibo radio telescope (Giovanelli et al. 2005). Data will be gathered with an homogeneous sensitivity (rms=2.5 mJy) and spectral resolution (5.5 km s^{-1}): at the average distance of the HRS the ALFALFA survey will detect all sources with $MHI \geq 10^{7.5} M_\odot$. VLA and WSRT HI maps are available for 236 HRS galaxies, from the WHISP (van der Hulst 2002) and VIVA (Chung et al. 2009) survey, this last limited to the Virgo region, from VLA archives or from our own WSRT observations.

A $^{12}\text{CO}(1-0)$ line (2.6 mm, 115 GHz) survey of the HRS galaxies is under way at the 12m Kitt Peak telescope. Fifty-eight spiral galaxies have been observed to date, with an average rms of 3 mK in T_R^* scale. Thanks to these new observations and to data in the literature CO measurements are now available for 161/258 late-type and 22/65 early-type galaxies with a detection rate of 83% and 55% respectively. Data have been taken from different sources, mostly from the FCRAO survey of Young et al. (1995), or Kenney & Young (1988) and Boselli et al. (1995; 2002) for the Virgo cluster region. We obtained time at the James Clerk Maxwell Telescope (JCMT) on the 345 GHz HARP array receiver to search for CO(3-2) emission from the central $3' \times 3'$ region (with 15 arcsec resolution) of a subsample of late type galaxies with $K_{\text{Stot}} < 8.7$. As of the beginning of January 2010, 42 of a total of 56 galaxies in the subsample have been observed to completion.

5.4. *Mid-IR, far-IR and sub-mm*

Infrared data for the HRS galaxies are already available from different sources. 208/258 late-type and 32/65 of the early-type galaxies have been detected by IRAS at 60 and 100 μm , only 103 (Sa-Sd-Im-BCD) and 17 (E-S0-S0a) at 12 and 25 μm . Virgo galaxies have been observed with CAM (45), PHOT (26) and LWS (21) on ISO (Leech et al. 1999; Malhotra et al. 2001; Roussel et al. 2001; Boselli et al. 2002, 2003; Tuffs et al. 2002). Spitzer observations have been completed for 157 HRS galaxies with IRAC and 181 galaxies with MIPS. The data for all of these galaxies will eventually be available from the Spitzer archives. The pipeline-processed IRAC data is sufficient for science analysis, but MIPS data has been reprocessed using the MIPS Data Analysis Tools (Gordon et al. 2005) along with additional processing software.

All HRS galaxies have been observed during the recently completed AKARI all sky survey in the 9, 20,

70, 90 and 160 μm bands which covered $\sim 95\%$ of the whole sky with an angular resolution comparable to that of SPIRE (~ 30 arcsec at 90 μm) and a sensitivity \sim three times better than IRAS. Pointed observations with AKARI and SCUBA are also available of a few ellipticals.

5.5. Radio continuum data

The HRS galaxies have been observed by the NVSS survey at 20 cm (1415 MHz) (Condon et al. 1998) with an angular resolution of 45 arcsec and a sensitivity of 2.5 mJy. 22/65 and 159/258 of the early- and late-type galaxies respectively have been detected at 20 cm. The data of the FIRST survey, at 5 arcsec resolution, are also available (Becker et al. 1995). For a fraction of galaxies data at 6.3 and 2.8 cm are also available from Niklas et al. (1995) and Vollmer et al. (2004).

5.6. X-ray data

X-ray data at 0.2-4 keV from the Einstein Observatory imaging instruments (IPC and HRI) are available for 31 early-type and 52 late-type galaxies from Fabbiano et al. (1992) and Shapley et al. (2001). Ten HRS galaxies have been detected by the ROSAT all sky survey (Voges et al. 1999), while Chandra observations of many of the early-type objects within the HRS are still under way (Gallo et al. 2008).

Table 5 shows the availability of ancillary data for the HRS galaxies.

6. HRS DATA PROCESSING, DATA PRODUCTS AND WEB SITE

The HRS galaxies are distributed widely across the sky, the survey will progress along with the Herschel mission, that is expected to last for three and a half years, with the first six months for commissioning, performance verification, science demonstration and the last three years for science. Data for each single observation will be proprietary for one year after it is taken. During this time, the data will be processed using the most up-to-date pipeline developed by the SPIRE-Instrument Control Centre since some members are also part of our consortium. This ensures we have access to the latest calibration files and bug fixes. Aside from the default data processing, the pipeline will also include new steps developed by the key-project teams (such as cosmic ray removal and baseline subtraction) for producing the most accurate representation of large-scale structures, such as the extended emission around galaxies. Additional post-processing steps will be applied to prepare the data so that it can be easily used by the broader astronomical community. This will include reformatting the headers of the fits files so that they are easy to understand. Furthermore, the data will have been visually inspected so as to ensure that they suffer from no severe artefacts related to the observations or data reduction, and in cases where problems are identified, the data will be manually re-processed to obtain the best results.

After the proprietary time the team plans to make the reduced data available to the community through a dedicated web page. Herschel and ancillary data will be archived in an Information System at the Laboratoire d'Astrophysique de Marseille developed under the SITools middleware interface

TABLE 3
THE COMPLETENESS IN THE DATA SAMPLE

λ	E+S0+S0a	Sa-Sm-Im-BCD
K	100%	100%
HI	85% (40%)	97% (93%)
FIR (60 μm)	95% (62%)	89% (83%)
Radio cont. (20 cm)	100% (34%)	100% (62%)

Note: Values in parenthesis give the detection rate.

(<http://vds.cnes.fr/sitools/>) with a Postgresql database. Images will be stored as fits files, and distributed in fits format while catalogues will be distributed in VO table and ASCII (cvs) format. The SPIRE data delivered by the team will be optimized for use by the wider astronomical community.

Given the legacy nature of this project, we will also present through an interactive database all the ancillary data, X, UV, optical, near- and far-IR, sub-mm and radio continuum fluxes and images, optical and line spectroscopy and derived parameters (structural parameters, metallicities, SFR...) with the aim of providing the community with a unique dataset, a suitable reference for future studies.

7. THE STATISTICAL PROPERTIES OF THE HERSCHEL REFERENCE SAMPLE

In this section, we investigate how representative the HRS is of the galaxy population as a whole, thus unbiased versus cosmic variance. This is necessary since, as selected, the HRS is biased versus galaxies located in a high density environment (Virgo). One way to test this is to see whether the luminosity distributions at different wavelengths in the HRS are similar to the local galaxy luminosity functions at the corresponding frequencies.

Shown in Fig. 3 is a comparison of the HRS luminosity distribution in the K-band (histogram) to the K band luminosity function of Cole et al. (2001) (dotted-dashed line; $\alpha = 0.96 \pm 0.05$, $M_K^* = -23.44 \pm 0.03$) and Kochanek et al. (2001) (solid line; $\alpha = 1.09 \pm 0.06$, $M_K^* = -23.39 \pm 0.05$), this last also considering separately early- (red dotted line; $\alpha = 0.92 \pm 0.10$, $M_K^* = -23.53 \pm 0.06$) and late-type (blue dashed line; $\alpha = 0.87 \pm 0.09$, $M_K^* = -22.98 \pm 0.06$) galaxies. The Kochanek et al. (2001) values of the K band luminosity function have been corrected to take into account the different assumed H_0 , while the Cole et al. (2001) M_K^* Kron magnitudes have been converted into total magnitudes as in this work. The zero point of the plotted luminosity functions is determined assuming the same number of objects as HRS within the given magnitude limit. Figure 3 shows that, when considered from the perspective of the K-band, the HRS is a good approximation to a volume-limited sample down to an absolute magnitude of $M_K \simeq -20$. This absolute magnitude limit is relatively high and explains why the HRS is under-represented in Im and BCD galaxies (Fig. 2). The sample includes only the brightest ellipticals and lenticulars, whose K band luminosity function ends at $M_K \sim -18$, while it does not include any quiescent dwarf system ($M_K \geq -20$).

Figure 4 shows the HI mass distribution of the HI detected HRS early-type (black histogram) and late-

type (gray histogram) galaxies compared to the HIPASS HI mass function determined by Zwaan et al. (2005) (dashed line, $\alpha = -1.37 \pm 0.03$, $M_{HI}^* = 10^{9.8 \pm 0.03} M_{\odot}$ and ϕ^* arbitrarily normalized to match the data). The sample includes galaxies with HI masses ranging in between $\sim 10^8$ and $10^{10} M_{\odot}$ but it is clearly under-representing galaxies with HI masses in the range $10^{7.5} M_{\odot} \leq M_{HI} \leq 10^9 M_{\odot}$ which are easily detectable by an HI blind survey such as ALFALFA. The HIPASS HI mass function is dominated by gas rich, star forming galaxies, while the HRS also includes quiescent, gas-poor early-types and HI-deficient cluster galaxies, as clearly observed in the Virgo cluster (Gavazzi et al. 2005; 2008) or in A1367 (Cortese et al. 2008b). The difference between the HIPASS mass function and the HRS HI mass distribution is due to i) the fact that the HRS is K band selected and thus does not include gas rich, low luminosity star forming systems with HI masses in the range $10^{7.5} \leq M_{HI} \leq 10^{8.5} M_{\odot}$ with K band magnitudes $12 \leq K \leq 16$, ii) the presence of cluster HI-deficient late-type galaxies. Indeed the HI mass distribution of galaxies with a normal HI gas content (HI-deficiency³⁰ < 0.4) is more similar to the HIPASS mass function than the distribution of gas poor objects (HI-deficiency ≥ 0.4). Most of the late-type HI-deficient objects (HI-deficiency ≥ 0.4) in the HRS are indeed located in the core of the Virgo cluster as shown in Fig. 5. Galaxies in the outskirts of the Virgo cluster or in the surrounding clouds have a normal HI gas content (see Table 4) and can be generally considered as unperturbed objects (unless belonging to groups or pairs) in the study of the effects of the environment on the dust properties of galaxies. Indeed, despite the relatively poor statistics, the average HI-deficiency of late-type galaxies in the Coma I Cloud, Leo Cloud, Ursa Major Cloud and Southern Spur, Crater Cloud, Canes Venatici Spur, Canes Venatici - Camelopardalis Cloud and Virgo - Libra Cloud is consistent with that of unperturbed galaxies in the reference sample of Haynes & Giovanelli (1984; HI-deficiency ≤ 0.3). The HI-deficiency of late-type galaxies in the different substructures of the Virgo cluster are generally consistent with those observed in a larger sample by Gavazzi et al. (1999a). The only exception is the Virgo E Cloud that we found dominated by HI-deficient objects. We notice however that the majority of the most HI-deficient galaxies (HI-deficiency > 1.0) of the HRS belonging to this substructure are classified as early-types in the VCC.

The Coma I Cloud, previously thought to include gas deficient objects (Garcia-Barreto et al. 1994; Gerin & Casoli 1994), is composed of galaxies with a normal gas content (Boselli & Gavazzi 2009). A detailed and complete study of the HI properties of late-type galaxies belonging to the other Clouds has not been done so far. A rather normal HI gas content of the spiral galaxies in the Ursa Major Cluster, a substructure of the Ursa Major Cloud, has been observed by Verheijen & Sancisi (2001).

³⁰ The HI-deficiency parameter is defined as the logarithmic difference between the average HI mass of a reference sample of isolated galaxies of similar type and linear dimension and the HI mass actually observed in individual objects: $HI - def = \text{Log}MHI_{ref} - \text{Log}MHI_{obs}$. According to Haynes & Giovanelli (1984), $\text{Log}MHI_{ref} = a + b \times \text{Log}(diam)$, where a and b are weak functions of the Hubble type and $diam$ is the linear diameter of the galaxy (see Gavazzi et al. 2005).

Figure 6 shows the far-IR luminosity distribution of the HRS early-type (black histogram) and late-type (gray histogram) galaxies detected by IRAS at $60 \mu\text{m}$ compared the IRAS $60 \mu\text{m}$ luminosity function determined by Takeuchi et al. (2003) ($\alpha = 1.23 \pm 0.04$, $L^* = 8.85 \times 10^8 L_{\odot}$ and ϕ^* in arbitrary units). The luminosity distribution of the HRS follows the luminosity function over two orders of magnitude. The low-luminosity cutoff is due to the detection limit of the IRAS survey. The lack of very luminous galaxies is because Luminous (LIRGs, $L_{60\mu\text{m}} = 10^{11} L_{\odot}$) and Ultra Luminous (ULIRGs, $L_{60\mu\text{m}} = 10^{12} L_{\odot}$) Infrared Galaxies are quite rare in the nearby universe and in particular inside rich clusters (Boselli & Gavazzi 2006).

Figure 7 shows the radio continuum (1.4 GHz) luminosity distribution of the radio detected HRS early-type (black histogram) and late-type (gray histogram) galaxies, as compared to the 2dF/NVSS luminosity function of Mauch & Sadler (2007) (solid line). This radio luminosity function includes the contribution from both AGN (dashed line) and quiescent (dotted line) galaxies. The radio luminosity distribution of the HRS is quite typical of normal, nearby galaxies ($10^{20} \leq L_{Radio} \leq 10^{22} \text{ W Hz}^{-1}$; Condon et al. 2002). The two brightest sources are the radio galaxies M87 (Virgo A, $L_{Radio} = 10^{24.8} \text{ W Hz}^{-1}$) and M84 ($L_{Radio} = 10^{23.3} \text{ W Hz}^{-1}$). The radio continuum luminosity distribution as a whole agrees well with the field luminosity function, indicating that the contamination of cluster galaxies with an enhanced radio continuum emission is minimal (Gavazzi & Boselli 1999c,d).

In summary, since the K-band light traces stellar mass, the good agreement between the HRS luminosity distribution and K-band luminosity function implies that the HRS can be treated as a volume-limited sample down to a minimum stellar mass. Figure 6 shows that the HRS is also representative of a far-IR selected sample once the most active and rare in the local universe LIRG and ULIRG are excluded. The HRS matches also the distribution of a radio continuum selected sample, although the statistic for the brightest radio galaxies is poor. However, Fig. 4 shows that when viewed from the perspective of gas mass, the HRS is not a fair sample of the local universe. The presence of HI-deficient objects makes it ideal for studying the effects of the cluster environment on the dust properties of nearby galaxies.

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TABLE 4
THE AVERAGE HI-DEFICIENCY OF LATE-TYPE GALAXIES IN THE DIFFERENT SUBSTRUCTURES

Substructure	average	standard deviation	N. of objects
Virgo A	0.91	± 0.47	33
Virgo B	0.71	± 0.46	25
Virgo N Cloud	0.39	± 0.24	12
Virgo E Cloud	0.83	± 0.53	12
Virgo S Cloud	0.47	± 0.47	26
Virgo Outskirts	0.26	± 0.60	38
Coma I Cloud	0.07	± 0.35	8
Leo Cloud	0.11	± 0.31	45
Ursa Major Cloud	0.05	± 0.34	15
Ursa Major Southern Spur	-0.02	± 0.20	4
Crater Cloud	0.13	± 0.40	4
Canes Ven. Spur & Camelopardalis	0.19	± 0.03	9
Virgo - Libra Cloud	0.31	± 0.35	18

Note: Upper limits are considered as detections

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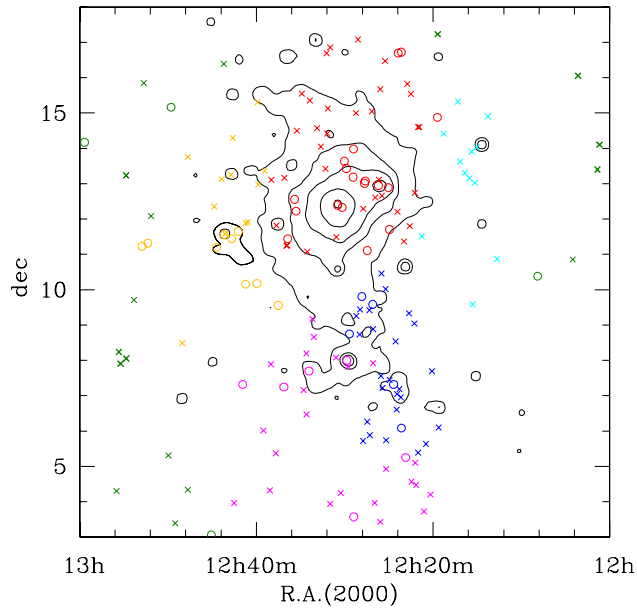
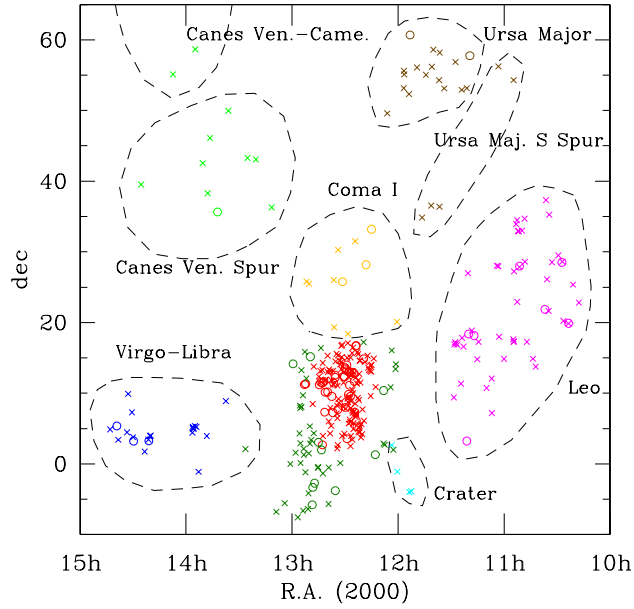


FIG. 1.— The sky distribution of the HRS for early-type (E-S0-S0a; circles) and late-type (Sa-Sd-Im-BCD; crosses) galaxies (left panel). Dashed contours delimit the different clouds. The big concentration of galaxies in the centre of the figure is the Virgo cluster (red colour) with its outskirts (dark green). Orange symbols are for Coma I Cloud, magenta for Leo Cloud, brown for Ursa Major Southern Spur and Cloud, cyan for Crater Cloud, light green for Canes Venatici Spur and Camelopardalis and blue for Virgo-Lybra Cloud galaxies respectively. The Virgo cluster region (right panel): black contours show the diffuse X-ray emission of the cluster (from Böhringer et al. 1994). Red symbols are for galaxies belonging to the Virgo A cloud, blue to Virgo B, orange to Virgo E, magenta to Virgo S, cyan to Virgo N and dark green to the Virgo outskirts, as defined by Gavazzi et al. (1999a).

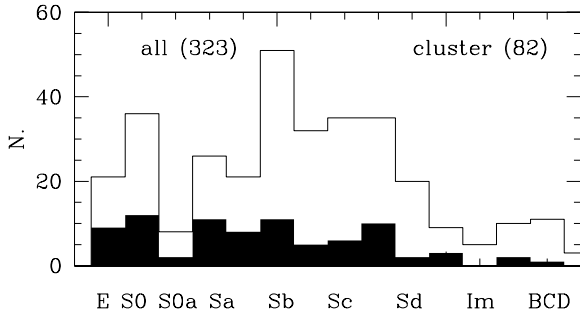


FIG. 2.— The distribution in morphological type of the HRS for all (empty histogram) and cluster (filled histogram) galaxies. The cluster sample is composed of galaxies members of the Virgo A and B clouds.

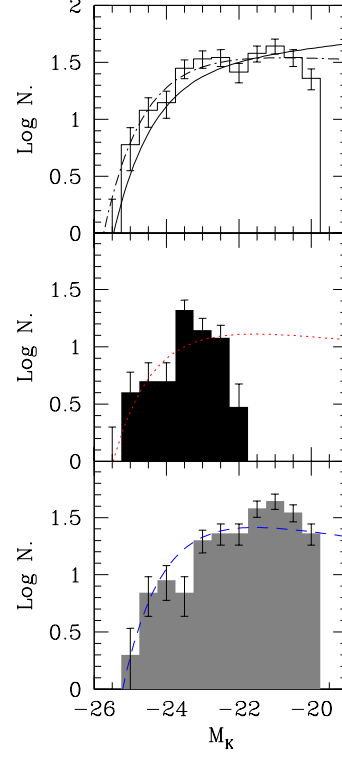


FIG. 3.— The K band luminosity distribution of the HRS for all galaxies is compared to the 2MASS K band luminosity function of Kochanek et al. (2001) (black solid line) and Cole et al. (2001) (black dotted-dashed line)(upper panel). The K band luminosity distributions of the E-S0-S0a (central panel) and Sa-Sd-Im-BCD (lower panel) galaxies in the HRS are compared to the Kochanek et al. (2001) K band luminosity function of early-type (red dotted line) and late-type (blue dashed line) galaxies. Poisson errors are also indicated.

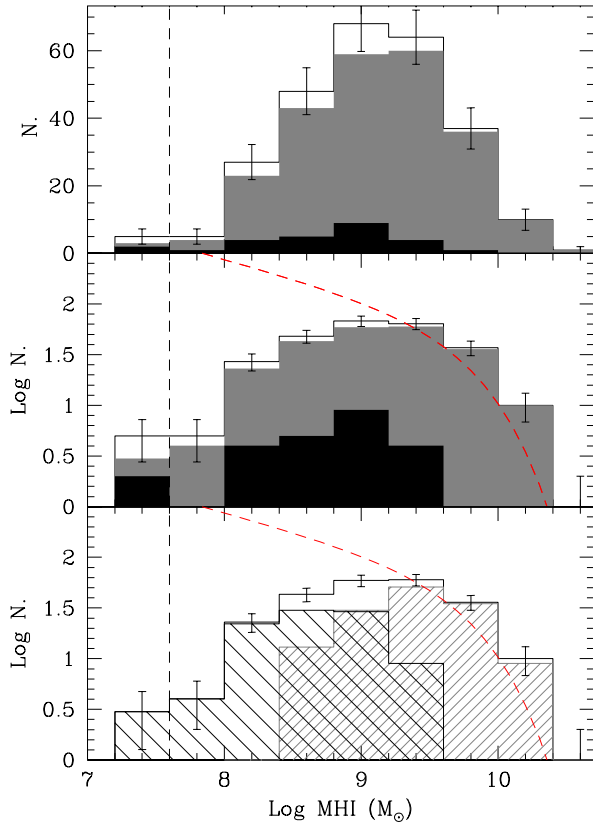


FIG. 4.— The atomic hydrogen mass distribution in linear (upper panel) and logarithmic (middle panel) scales for the HI detected HRS galaxies (empty histogram). The black and gray histograms are for early-type (E-S0-S0a, 35% detected) and late-type (Sa-Sd-Im-BCD; 93 % detected) galaxies respectively. Lower panel: The atomic hydrogen mass distribution for late-type galaxies with a normal HI gas content (HI-deficiency < 0.4 ; tiny spaced hashed histogram) and gas poor galaxies (HI-deficiency ≥ 0.4 ; large spaced hashed histogram). The red dashed line is the HI mass function determined by Zwaan et al. (2005). The vertical, dashed line indicates the detection limit of the ALFALFA survey (2.5 mJy) at the distance of 20 Mpc.

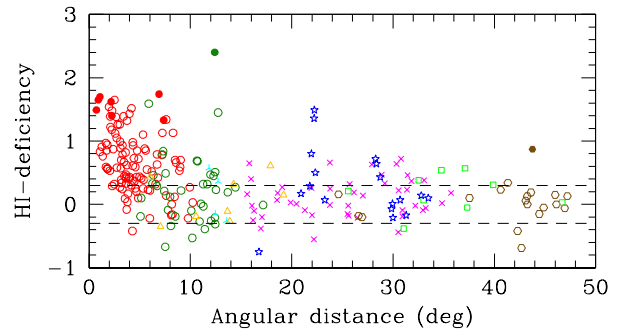


FIG. 5.— The variation of the HI-deficiency parameter as a function of the angular distance from the core of the Virgo cluster (M87). Red circles are for Virgo cluster galaxies (Virgo A, B, N, S and E), dark green circles for galaxies in the Virgo outskirts, orange triangles for the Coma I Cloud, magenta crosses for the Leo Cloud, brown hexagons for Ursa Major Southern Spur and Cloud, cyan three skeletal for the Crater Cloud, light green squares for Canes Venatici Spur and Camelopardalis and blue stars for the Virgo-Lybra Cloud. Empty symbols are for HI-detected galaxies, filled symbols for upper limits.

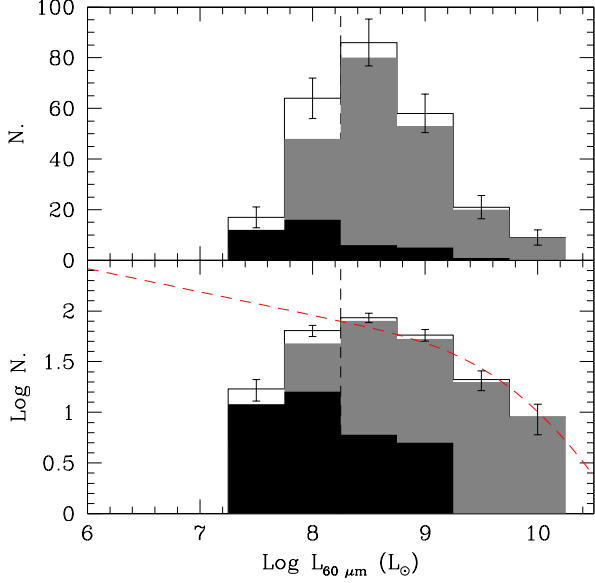


FIG. 6.— The infrared luminosity distribution in linear (upper panel) and logarithmic (lower panel) scales for the FIR ($60 \mu\text{m}$) detected HRS galaxies (empty histogram). The black and gray histograms are for early-type (E-S0-S0a) and late-type (Sa-Sd-Im-BCD) galaxies respectively. The dashed line is the IRAS $60 \mu\text{m}$ luminosity function determined by Takeuchi et al. (2003). The vertical, dashed line indicates the typical detection limit of IRAS (0.4 mJy) at the distance of 20 Mpc.

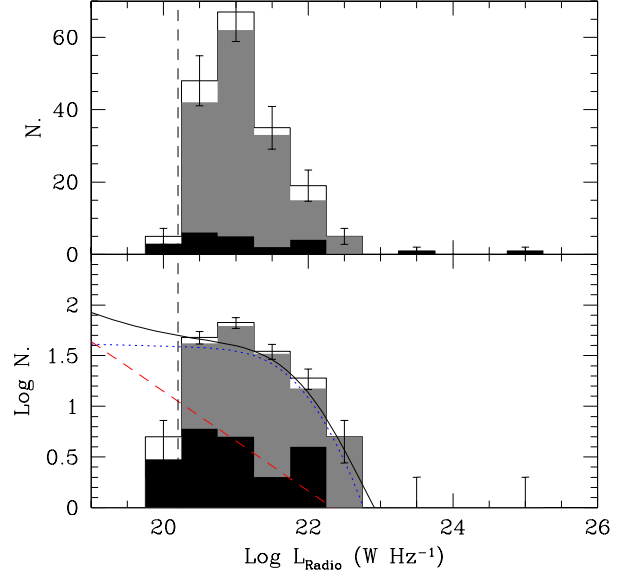


FIG. 7.— The radio continuum (1415 MHz) luminosity distribution in linear (upper panel) and logarithmic (lower panel) scales for the detected HRS galaxies (empty histogram). The black and gray histograms are for early-type (E-S0-S0a) and late-type (Sa-Sd-Im-BCD) galaxies respectively compared to the 2dF/NVSS luminosity function of Mauch & Sadler (2007) (solid line). This radio luminosity function includes both the contribution of AGN (red dashed line) and quiescent (blue dotted line) galaxies. The vertical, dashed line indicates the typical detection limit of NVSS (2.5 mJy) at the distance of 20 Mpc.

TABLE 5
 ANCILLARY DATA FOR THE HERSCHEL REFERENCE SAMPLE.

HRS	2MASS	SDSS	UV	H α	IRAS	20cm	HI	CO	Spec
1	X	X	-	-	X	-	-	-	X
2	X	X	X	-	X	X	X	-	X
3	X	X	-	X	-	X	X	X	-
4	X	X	-	X	X	X	X	X	X
5	X	X	-	X	X	X	X	-	X
6	X	X	X	-	-	-	X	-	-
7	X	X	X	-	X	X	X	X	-
8	X	X	X	X	X	-	X	X	X
9	X	X	X	X	X	X	X	X	X
10	X	X	-	-	X	X	X	-	X
11	X	X	X	X	X	X	X	X	X
12	X	X	X	-	X	-	X	-	X
13	X	X	X	X	X	X	X	X	X
14	X	X	X	-	X	X	X	-	-
15	X	X	X	X	X	X	X	X	X
16	X	X	X	X	X	-	X	X	X
17	X	X	X	X	X	X	X	X	X
18	X	X	X	X	X	X	X	-	X
19	X	X	X	X	X	X	X	-	X
20	X	X	X	X	-	X	X	X	X
21	X	X	X	X	-	-	X	-	X
22	X	X	X	X	-	X	X	-	-
23	X	X	X	X	X	X	X	X	X
24	X	X	X	X	X	X	X	X	X
25	X	X	X	X	X	X	X	X	X
26	X	X	X	-	X	-	X	-	X
27	X	X	X	X	X	X	X	-	X
28	X	X	X	X	X	X	X	-	X
29	X	X	X	-	-	-	X	-	X
30	X	X	X	-	-	X	X	-	X
31	X	X	X	X	X	X	X	X	X
32	X	X	X	X	-	-	X	X	X
33	X	X	X	X	-	X	X	X	X
34	X	X	X	X	-	X	X	X	X
35	X	X	-	-	-	-	-	-	X
36	X	X	X	X	X	X	X	X	X
37	X	X	X	X	X	X	X	X	X
38	X	X	-	X	X	X	X	X	X
39	X	X	X	-	-	-	X	-	X
40	X	X	X	X	-	X	X	-	X
41	X	X	X	X	-	-	X	-	-
42	X	X	-	X	-	X	X	X	X
43	X	X	X	-	-	-	X	-	-
44	X	X	X	-	-	X	X	-	X
45	X	X	X	-	X	X	X	-	-
46	X	X	-	-	-	X	X	-	-
47	X	X	X	X	-	X	X	-	X
48	X	X	X	X	X	X	X	X	X
49	X	X	X	-	-	-	X	-	-
50	X	X	X	X	X	X	X	X	X
51	X	X	-	X	X	X	X	-	X
52	X	X	X	-	X	X	X	-	X
53	X	X	X	X	X	X	X	X	X
54	X	X	X	X	X	-	X	X	X
55	X	X	X	X	X	X	X	X	X
56	X	X	X	X	X	X	X	X	X
57	X	X	X	X	X	X	X	X	X
58	X	X	X	X	X	X	X	-	X
59	X	X	X	X	X	X	X	X	X
60	X	X	X	X	X	X	X	X	X
61	X	X	X	-	X	-	X	-	X
62	X	X	X	X	-	X	X	-	X
63	X	X	-	X	X	-	X	X	X
64	X	X	-	-	X	-	X	-	X
65	X	X	X	-	X	-	X	-	X
66	X	X	X	X	-	X	X	X	X
67	X	X	-	-	-	X	X	-	X
68	X	X	X	-	-	X	X	-	X
69	X	X	-	X	X	X	X	X	X
70	X	-	X	X	X	X	X	-	X
71	X	X	X	-	X	-	-	-	-
72	X	-	X	-	X	X	X	-	X
73	X	X	-	X	X	-	X	X	X
74	X	X	-	X	X	X	X	X	X
75	X	X	-	-	-	-	X	-	X
76	X	X	-	X	X	-	X	-	X
77	X	X	X	X	X	X	X	X	X
78	X	X	X	X	X	-	X	X	X
79	X	X	X	X	X	X	X	-	X
80	X	X	X	X	X	-	X	-	X
81	X	X	X	X	X	X	X	X	X
82	X	X	X	-	X	X	X	-	X
83	X	X	X	-	X	-	X	-	X
84	X	X	X	X	X	X	X	X	X
85	X	X	X	X	X	X	X	X	X
86	X	X	X	X	X	X	X	X	X
87	X	X	X	-	X	-	X	-	X
88	X	X	X	X	X	X	X	X	X
89	X	X	X	X	X	X	X	X	X
90	X	X	X	-	-	-	X	-	X

TABLE 5
CONTINUE

HRS	2MASS	SDSS	UV	H α	IRAS	20cm	HI	CO	Spec
91	X	X	-	X	X	X	X	X	X
92	X	X	-	X	X	-	X	-	X
93	X	X	X	-	X	X	X	X	-
94	X	X	X	X	X	-	X	X	X
95	X	X	X	X	X	X	X	X	X
96	X	X	-	X	X	X	X	X	X
97	X	X	X	X	X	X	X	X	X
98	X	X	X	X	X	X	X	X	X
99	X	X	X	X	X	-	X	-	X
100	X	X	X	X	X	X	X	X	X
101	X	X	X	-	X	-	X	X	-
102	X	X	X	X	X	X	X	X	X
103	X	X	X	X	X	-	X	X	X
104	X	X	X	X	-	-	-	-	-
105	X	X	-	-	-	-	X	-	X
106	X	X	X	X	X	-	X	-	X
107	X	X	X	X	-	-	X	-	X
108	X	X	X	X	-	-	X	-	X
109	X	X	X	X	X	X	X	-	X
110	X	X	X	X	X	X	X	X	X
111	X	X	X	X	X	X	X	X	X
112	X	X	-	X	X	-	X	X	X
113	X	X	X	X	X	X	X	X	X
114	X	X	X	X	X	X	X	X	X
115	X	X	X	X	-	-	X	-	X
116	X	X	-	X	-	-	X	-	X
117	X	X	-	X	X	X	X	X	X
118	X	X	X	X	X	-	X	-	X
119	X	X	X	X	X	X	X	X	X
120	X	X	X	X	X	-	X	X	X
121	X	X	X	X	X	X	X	X	X
122	X	X	X	X	X	X	X	X	X
123	X	X	X	X	X	-	X	-	X
124	X	X	X	X	X	X	X	X	X
125	X	X	X	-	-	-	X	-	-
126	X	X	-	-	X	-	X	-	-
127	X	X	X	X	X	X	X	X	X
128	X	X	X	X	X	-	X	-	X
129	X	X	X	-	X	-	X	-	X
130	X	X	X	X	X	-	X	X	X
131	X	X	X	X	-	-	X	-	X
132	X	X	X	X	X	X	X	-	X
133	X	X	X	X	X	-	X	X	X
134	X	X	-	X	X	-	X	-	X
135	X	X	X	X	-	-	-	-	X
136	X	X	-	X	X	-	X	X	X
137	X	X	X	-	-	-	X	-	X
138	X	X	X	X	X	X	X	X	-
139	X	X	X	X	X	-	X	-	X
140	X	X	X	X	X	-	X	X	X
141	X	X	X	X	X	-	X	X	X
142	X	X	X	X	X	X	X	X	X
143	X	X	X	X	X	X	X	-	X
144	X	X	X	X	X	X	X	X	X
145	X	X	X	X	X	-	X	X	X
146	X	X	X	X	X	X	X	-	X
147	X	X	X	X	X	-	X	-	X
148	X	X	X	X	X	X	X	X	X
149	X	X	X	X	X	X	X	X	X
150	X	X	X	X	X	-	-	X	X
151	X	X	X	X	X	-	X	X	X
152	X	X	X	X	X	X	X	X	X
153	X	X	X	X	X	X	X	X	X
154	X	X	X	X	X	-	X	X	X
155	X	X	X	-	-	-	X	-	X
156	X	X	X	X	X	X	X	X	X
157	X	X	X	X	X	-	X	X	X
158	X	X	X	X	X	X	X	X	X
159	X	X	X	X	X	X	X	X	X
160	X	X	X	X	X	X	X	X	X
161	X	X	X	X	X	-	X	X	X
162	X	X	X	-	X	-	X	X	X
163	X	X	X	X	X	X	X	X	X
164	X	X	X	X	-	-	X	X	X
165	X	X	X	X	-	-	X	-	X
166	X	X	X	-	-	-	X	-	X
167	X	X	X	X	X	-	X	-	X
168	X	X	X	X	X	X	X	-	X
169	X	X	X	X	X	-	X	-	X
170	X	X	X	X	X	X	X	X	X
171	X	X	X	X	X	X	X	X	X
172	X	X	X	X	X	X	X	X	X
173	X	X	X	X	X	X	X	X	X
174	X	X	X	X	X	-	X	X	X
175	X	X	X	X	-	-	X	-	X
176	X	X	X	X	X	-	X	-	X
177	X	X	X	X	X	X	X	X	X
178	X	X	X	X	-	X	X	X	X
179	X	X	X	X	-	-	-	-	X
180	X	X	X	-	X	-	X	-	X

TABLE 5
CONTINUE

HRS	2MASS	SDSS	UV	H α	IRAS	20cm	HI	CO	Spec
181	X	X	X	-	-	-	-	-	-
182	X	X	X	X	X	X	X	X	X
183	X	X	X	X	-	X	-	X	X
184	X	X	X	X	X	-	X	X	X
185	X	X	X	X	-	-	X	X	X
186	X	X	-	-	-	-	X	X	-
187	X	X	X	X	X	-	X	X	X
188	X	X	X	X	X	X	X	X	X
189	X	X	X	X	X	X	X	-	X
190	X	X	X	X	X	X	X	X	X
191	X	X	X	X	X	-	X	-	X
192	X	X	X	X	X	-	X	-	X
193	X	X	X	X	X	X	X	X	X
194	X	X	X	X	X	-	X	X	X
195	X	X	-	X	-	-	X	X	X
196	X	X	X	X	X	X	X	X	X
197	X	X	X	X	X	X	X	X	X
198	X	X	X	X	X	-	X	X	X
199	X	X	X	X	X	-	X	-	X
200	X	X	X	X	X	X	X	X	-
201	X	X	X	X	X	X	X	X	X
202	X	X	-	-	-	-	-	-	X
203	X	X	X	X	X	X	X	X	X
204	X	X	X	X	X	X	X	X	X
205	X	X	X	X	X	X	X	X	X
206	X	X	X	X	X	X	X	X	X
207	X	X	X	X	X	X	X	X	X
208	X	X	X	X	X	-	X	X	X
209	X	-	X	-	X	X	X	-	-
210	X	X	X	-	-	-	X	-	X
211	X	X	X	X	X	X	-	-	X
212	X	X	X	X	X	X	X	X	X
213	X	X	X	X	X	X	X	X	X
214	X	X	X	-	-	-	X	-	-
215	X	X	X	X	-	X	X	X	X
216	X	X	X	X	X	X	X	X	X
217	X	X	X	X	X	X	X	X	X
218	X	X	X	-	-	-	X	-	-
219	X	X	-	-	-	-	X	-	-
220	X	X	X	X	X	X	X	X	X
221	X	X	X	X	X	-	X	X	X
222	X	X	X	X	-	-	X	-	X
223	X	X	X	X	-	-	X	-	X
224	X	X	X	X	X	-	X	X	X
225	X	X	X	X	-	-	X	-	X
226	X	X	X	X	X	-	X	-	X
227	X	X	X	X	X	X	X	-	X
228	X	-	-	-	-	-	-	-	-
229	X	X	X	-	-	-	X	-	X
230	X	X	X	X	X	X	X	X	X
231	X	X	-	X	X	-	X	X	X
232	X	X	X	X	X	-	X	X	X
233	X	X	X	X	X	X	X	X	X
234	X	X	-	-	-	X	X	-	X
235	X	X	X	-	-	-	X	-	-
236	X	X	X	-	-	-	-	-	X
237	X	X	X	X	X	X	X	X	X
238	X	X	-	-	-	-	X	-	X
239	X	X	X	X	X	X	X	X	X
240	X	X	X	-	-	-	X	-	X
241	X	X	X	X	-	X	X	X	-
242	X	X	X	X	X	X	X	X	X
243	X	X	X	X	X	-	X	-	-
244	X	X	X	X	X	X	X	X	X
245	X	X	X	X	X	X	X	X	X
246	X	X	X	X	X	X	X	X	X
247	X	X	X	X	X	X	X	X	X
248	X	X	X	-	-	-	-	-	X
249	X	X	X	X	-	-	X	-	X
250	X	X	X	-	-	-	X	-	-
251	X	X	X	X	X	X	X	X	X
252	X	X	X	-	X	-	X	-	X
253	X	X	X	-	X	X	X	-	-
254	X	X	-	X	X	X	X	X	X
255	X	X	X	X	X	X	X	X	X
256	X	X	X	-	X	X	X	X	-
257	X	X	X	X	X	-	X	X	X
258	X	X	X	X	X	-	-	X	-
259	X	X	-	X	X	X	X	X	X
260	X	X	X	X	X	X	X	X	X
261	X	X	X	-	X	-	X	-	X
262	X	X	X	X	X	X	X	X	X
263	X	X	X	X	X	X	X	X	X
264	X	X	X	-	-	-	X	-	X
265	X	-	X	-	X	X	-	-	X
266	X	-	X	X	X	X	X	X	X
267	X	X	X	X	X	X	X	-	X
268	X	X	X	X	X	X	X	X	X
269	X	X	X	-	-	-	X	-	X
270	X	X	X	-	X	-	X	X	X

